

A COMPARATIVE ASTROCHEMICAL STUDY OF THE HIGH-MASS PROTOSTELLAR OBJECTS NGC 7538 IRS 9 AND IRS 1

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ABSTRACT

We report the results of a spectroscopic study of the high-mass protostellar object NGC 7538 IRS 9 and compare our observations to published data on the nearby object NGC 7538 IRS 1. Both objects originated in the same molecular cloud and appear to be at different points in their evolutionary histories, offering an unusual opportunity to study the temporal evolution of envelope chemistry in objects sharing a presumably identical starting composition. Observations were made with the Texas Echelon Cross Echelle Spectrograph (TEXES), a sensitive, high spectral resolution ($R = \lambda/\Delta\lambda \simeq 100,000$) mid-infrared grating spectrometer. Forty-six individual lines in vibrational modes of the molecules C_2H_2 , CH_4 , HCN , NH_3 and CO were detected, including two isotopologues (^{13}CO , $^{12}C^{18}O$) and one combination mode ($\nu_4 + \nu_5$ C_2H_2). Fitting synthetic spectra to the data yielded the Doppler shift, excitation temperature, Doppler b parameter, column density and covering factor for each molecule observed; we also computed column density upper limits for lines and species not detected, such as $HNCO$ and OCS . We find differences among spectra of the two objects likely attributable to their differing radiation and thermal environments. Temperatures and column densities for the two objects are generally consistent, while the larger line widths toward IRS 9 result in less saturated lines than those toward IRS 1. Finally, we compute an upper limit on the size of the continuum-emitting region (~ 2000 AU) and use this constraint and our spectroscopy results to construct a schematic model of IRS 9.

Subject headings: stars: formation — stars: pre-main sequence — techniques: spectroscopic — ISM: individual objects (NGC 7538 IRS 9, NGC 7538 IRS 1)

1. INTRODUCTION

Details of the pre-Main Sequence (pre-MS) life cycles of low-mass stars such as the Sun ($M \sim 1M_\odot$) have been gleaned from decades of research; see McKee & Ostriker (2007) for a recent review. This is thanks in part to the relative amenability of low-mass protostellar objects to study. However, details of high-mass ($M > 8M_\odot$) star formation are not as well understood at present, in part because high-mass stars are not generally forming near the solar neighborhood and remain deeply embedded during their pre-MS evolution. An important unresolved issue is whether high-mass stars form in a manner resembling a scaled-up version of low-mass star formation (Keto & Zhang 2010; Johnston et al. 2011) or through other mechanisms uncharacteristic of the low-mass case such as collisions/mergers (Bonnell et al. 1998; Bonnell & Bate 2005; Bally & Zinnecker 2005) and competitive accretion in clusters (Bonnell et al. 2001, 2004). Clearly, the environments in which high-mass star formation takes place seem to influence the formation process itself. Massive stars are most often seen to form in large groups (“OB associations”) near the edges of dense clumps in giant molecular clouds (GMCs; Zinnecker & Yorke 2007). Elmegreen & Lada (1977) first suggested that spatially distinct subgroups of stars within OB associations trigger waves of self-propagating star formation in GMCs via shock/ionization fronts that progress from the outside in. This process may account for observations of OB subgroups that lie in linear arrangements along the Galactic plane, showing a monotonically-increasing sequence of ages (Blaauw 1964, 1991). Rapid improve-

ments in both observational capabilities and theoretical sophistication in recent years are advancing our knowledge of the circumstances of massive star formation and enabling these fundamental questions to be addressed.

The origin of these stars is of particular interest due to their role in the Galactic ecosystem. In addition to serving as the principal sites of heavy element nucleosynthesis, the violent end of massive stars in supernovae injects considerable mechanical energy into the interstellar medium (ISM) while enriching it with metals (Arnett 1996). The debris are reconstituted in molecular clouds from which new stars form, often accompanied by circumstellar disks (Jiang et al. 2008). These disks are evidently the source of accretion onto high-mass stars as they approach the MS in the Hertzsprung-Russell diagram and ignite hydrogen burning, then continue to accrete their final masses, moving parallel to the MS (Zinnecker & Yorke 2007). Relatively little is known about the structure and chemical composition of massive protostellar envelopes, and even less about how they evolve in time. Ready comparison with envelope models for low-mass objects is instructive but only of limited use, in part due to the radically different radiation environment in the high-mass case. The composition and structure of the envelope are of interest as they determine the final mass of the star through an accretion mechanism whose details are still unclear. Piecing together the components of protostellar envelopes is complicated by the observation that most objects are deeply embedded, particularly early in the formation process. Spectroscopy of molecules is useful in delineating physical structures at high extinction

where other methods are limited. Pure rotational transitions of molecules in the millimeter probe a sufficiently low energy regime but such observations are typically of low spatial resolution. Better resolution is achieved in the infrared where many ro-vibrational molecular lines are located.

We present high resolution, mid-infrared spectra of a number of molecules toward an embedded infrared source in the NGC 7538 starforming complex, NGC 7538 IRS 9, and compare the results to a previous study of the nearby source NGC 7538 IRS 1 (Knez et al. 2009). Both sources are presumed to harbor high-mass protostellar objects in their interiors. Their proximity to one another suggests that, having formed within the same molecular cloud, the starting chemical abundances for each object should have been substantially similar. A comparison of their current states provides insights about differences in their ages and/or divergence in their individual evolutionary histories.

Weak free-free emission ($\leq 60 \mu\text{Jy}$ at 3.6 cm; Sandell et al. 2005) is observed toward IRS 9. Further evidence for the comparatively primitive nature of the envelope of IRS 9 is given by the detection of various ices in its *Infrared Space Observatory* (ISO) spectrum (Whittet et al. 1996; Gibb et al. 2004). These ices are not seen toward IRS 1, which appears to be in a relatively advanced point in its pre-MS evolution. It is a considerably stronger source at centimeter wavelengths and shows evidence of both ionized and molecular components in a bipolar outflow originating in a disk seen nearly edge-on (Sandell et al. 2009; Qiu et al. 2011). IRS 1 also shows signs of a circumstellar disk and an emerging compact H II region; it is thought to be a forming late O star undergoing active accretion, suggesting a more advanced evolutionary state (Lacy & Jaffe 2007; Beuther et al. 2012). Sandell et al. (2009) further note that the system is heavily accreting ($\dot{M} \sim 2 \times 10^{-4} M_{\odot} \text{ yr}^{-1}$) and that the accretion may be episodic in nature. On the basis of GHz line observations, Surcis et al. (2011) conclude that the accretion appears to be the result of radial infall from a “torus” thousands of Astronomical Units (AU) in size, traced by H_2O and CH_3OH masers, instead of the Keplerian disk favored by Pestalozzi et al. (2004).

Until recently, estimates of the distance to NGC 7538 placed it at ~ 3.0 kpc (Heyer et al. 2001; Brunt 2003; Foster & Routledge 2003; Balog et al. 2004; Reid & Wilson 2005; Kameya 2006; Araya et al. 2007). However, we adopted the distance of Moscadelli et al. (2009), who find a value of $2.65^{+0.12}_{-0.11}$ kpc from trigonometric parallax measurements. The total luminosity of IRS 9 is about $3.5 \times 10^4 L_{\odot}$ (Sandell et al. 2005, corrected to the Moscadelli et al. distance) and it is a bright IRAS 12 μm source at ~ 60 Jy (Beichman et al. 1988). A zero-age Main Sequence (ZAMS) star with this luminosity would have a spectral type of B0.5, an effective temperature $T_{\text{eff}} = 26200$ K (Panagia 1973), and a mass of $\sim 13 M_{\odot}$ (Ekström et al. 2012). Boogert et al. (2004) suggest the existence of an inner, warm molecular emission region with a radius of ~ 70 AU; the dust temperature at this distance for an object of IRS 9’s luminosity would be ~ 1000 K depending on the dust mass opacity. Millimeter data show that IRS 9 is a site of active, massive star formation (Sandell et al. 2005) in which young stellar objects

(YSOs) are driving a set of bipolar molecular outflows on a dynamical timescale of $\leq 20,000$ years. At least one of these outflows appears to be highly energetic. Mitchell & Hasegawa (1991) report the discovery of an outflow in IRS 9 with a velocity of 110 km s^{-1} , which they interpret to have emerged as recently as 1,200 years ago, perhaps indicating the beginning of a recent episode of accretion.

Observations of these protostellar objects are broadly consistent with the Elmegreen & Lada (1977) picture of sequential high-mass star formation, as elaborated on by others in the particular case of NGC 7538 (Werner et al. 1979; Fischer et al. 1980; Dickel et al. 1981). The Campbell & Thompson (1984) model of NGC 7538 holds that a first generation of massive stars caused the visible H II region, which expanded and impinged on the neighboring molecular cloud; this compressed the cloud and generated shocks that caused the observed near-infrared fluorescent H_2 emission northwest of IRS 1. In this picture, high-mass star formation is progressing toward the southeast, and IRS 9 belongs to a second wave of star formation as it is placed further along the direction of propagation of the shock front. These objects then represent two snapshots of the evolution of a massive young protostellar object in a particular cloud and are useful in addressing temporal evolution questions.

This paper is organized as follows. In Section 2 we give details concerning how we obtained and reduced the data, followed by a detailed description in Section 3 of the molecular species detected. The data analysis is described in Section 4 with emphasis on spectral modeling, and we give some interpretation to the results in Section 5. We summarize these results and draw some conclusions in Section 6.

2. OBSERVATIONS AND REDUCTIONS

We observed NGC 7538 IRS 9 with the Texas Echelon Cross (X) Echelle Spectrograph (TEXES; Lacy et al. 2002) on the Gemini North 8 m telescope in 2007 October. Similar observations of IRS 1 were carried out with TEXES at the NASA Infrared Telescope Facility (IRTF) 3 m telescope in 2001 June, 2001 November, 2002 September, 2002 December and 2005 November. Circumstances for all IRS 9 observations are given in Table 1. IRTF observations of IRS 1 form the basis for the work of Knez et al. (2009). Additional observations of IRS 1 were made at Gemini in 2007 October to supplement previous TEXES data taken at the IRTF.

TEXES is a cross-dispersed grating spectrograph designed to operate at high spectral resolution ($R = 75,000$ – $100,000$) over a range of mid-infrared wavelengths from 5 to 25 μm . Light diffracts off a coarsely-ruled “echelon” grating (Michelson 1898) at a large incidence angle with respect to normal at order $n \sim 1500$. The light then diffracts off a cross-dispersion grating and is imaged onto a 256×256 pixel Si:As detector as a series of orders. At a given spectral setting between five and ten cross-dispersed orders are imaged onto the detector for a spectral coverage of about 0.5%. The continuity of orders is wavelength-dependent; below 11 μm order overlap is sufficient to yield continuous coverage over an entire spectral setting. At wavelengths longer than 11 μm there are gaps between adjacent orders.

On Gemini the pixel scale of TEXES was $\sim 0''.14$ while at the IRTF it was $\sim 0''.36$. The pixel scales also compare

TABLE 1
 CIRCUMSTANCES OF TEXES OBSERVATIONS OF NGC 7538 IRS 9

UT Date	Telescope	Wavenumber Center (cm ⁻¹)	Included Features
2001 June 15	IRTF	768 780	C ₂ H ₂ ν_5 $R(15,16)$ [Ne II] ($^2P_{1/2} \rightarrow ^2P_{3/2}$), C ₂ H ₂ ν_5 $R(21,22)$, HCN ν_2 $R(22,23)$
2001 June 18	IRTF	744	C ₂ H ₂ ν_5 $R(5,6)$, HCN ν_2 $R(10)$
2001 June 26	IRTF	761	C ₂ H ₂ ν_5 $R(12,13)$, HCN ν_2 $R(15,16)$, HD, H ¹³ CN ν_2 $R(18)$
2001 June 28-29	IRTF	734	C ₂ H ₂ ν_5 $R(1)$, HCN ν_2 $R(6,7)$
2007 October 21	Gemini	828 1308	HNCO ν_4 P -branch, ¹⁴ NH ₃ ν_2 $sP(7,K)$ ¹² CH ₄ ν_4 $R(0)$, ¹³ CH ₄ ν_4 $R(1)$
2007 October 25	Gemini	930 743	¹⁴ NH ₃ ν_2 $aQ(J,K)$ C ₂ H ₂ ν_5 $R(5,6)$
2007 October 26	Gemini	853 1323	¹⁴ NH ₃ ν_2 $aP(4,K)$ ¹² CH ₄ ν_4 $R(2)$, ¹³ CH ₄ ν_4 $R(4)$, C ₂ H ₂ $\nu_4 + \nu_5$ $P(1-3)$
2007 October 27	Gemini	2055	OCS ν_2 $P(12-22)$, ¹² CO $v=1-0$ $P(21,22)$, ¹³ CO $P(11,12)$, C ¹⁸ O $P(10,11)$, C ¹⁷ O $P(16)$
2007 October 28	Gemini	730	C ₂ H ₂ ν_5 Q -branch
2007 October 30	Gemini	2085	¹² CO $v=1-0$ $P(14-15)$, ¹³ CO $P(3,4)$, C ¹⁸ O $P(2,3)$, C ¹⁷ O $P(8,9)$

favorably with the diffraction limit at 10 μm of each telescope, 0".70 and 0".26 at the IRTF and Gemini, respectively. In both cases the seeing disc is, at worst, critically sampled under good seeing conditions. All observations were made at a resolving power $R \simeq 80,000$, or $\Delta v \sim 3\text{--}4 \text{ km s}^{-1}$ and the spectral coverage at each spectral setting was $\Delta\lambda \approx 0.06 \mu\text{m}$. The spectral sampling of the TEXES array is about 1.0 km s^{-1} . Slit widths were 0".5 at Gemini and 1".4 at the IRTF. The slit length was $\sim 4''$ on Gemini and $\sim 10''$ on the IRTF, varying with wavelength. We alternated the telescope pointing between two positions on the sky in the slit direction, separated by $\sim 1\text{--}5''$, during an exposure sequence. Subtraction of adjacent "nod" positions removes the contribution to the signal from night sky emission. The separation of points in the nod, or "nod throw", is determined by the slit length.

To correct for telluric atmospheric absorption we observed comparison objects whose spectra are close to blackbody over the wavelength range of observation. At wavelengths beyond $\sim 8 \mu\text{m}$, asteroids meet this requirement while providing the most flux of available object types. Toward $5 \mu\text{m}$, most asteroids provide insufficient flux to achieve a high signal-to-noise (S/N) ratio, so we used bright, hot Main Sequence stars such as Vega (α Lyr), Fomalhaut (α PsA) and Mirfak (α Per) as divisors. We aimed to observe the divisor for a given target at nearly the same airmass as the target observation for better divisions. Differences in airmass at the times of observations may be corrected after data reduction by taking the divisor spectrum to the power of the ratio of the airmasses. Residual structure in divided spectra was removed by fitting and subtracting low-order polynomials.

Finally, given the high spatial resolution of the Gemini observations, we considered the possibility that TEXES might resolve the emitting/ absorbing region around IRS 9 if it were spatially extended. We obtained a series of spatial-spectral scan maps of IRS 9 that allow us to place an upper limit on the linear size of any resolved struc-

ture. The approach and results are discussed in detail in Section 5.5.

2.1. Data Reduction

Our data reduction pipeline is described in detail in Lacy et al. (2002). Flux calibration is based on the approach of calibrating millimeter and submillimeter data in the manner of Ulich & Haas (1976). A typical sky calibration cycle involves collecting data on the sky and object with two positions of a rotating chopper blade mounted just above the entrance window to the TEXES Dewar: one painted flat black and one low-emissivity, polished aluminum ("shiny") surface. The result of dividing by the black-sky difference removes atmospheric absorption and serves as both a spatial and spectral flat field image provided that the black position, sky, and telescope temperatures are approximately equal. We quote the Lacy et al. (2002) expression for the object intensity as a function of frequency,

$$I_\nu(\text{object}) \approx S_\nu(\text{object} - \text{sky}) \frac{B_\nu(T_{\text{tel}})}{S_\nu(\text{black} - \text{sky})} \quad (1)$$

where S_ν is the measured signal and B_ν is the blackbody function at the telescope temperature T_{tel} . In the approximation that $I_\nu(\text{black}) - I_\nu(\text{sky}) = B_\nu(T_{\text{tel}})e^{-\tau_{\text{sky}}}$, this procedure corrects for sky absorption. Wavelength calibration is achieved by using the known wavelengths of night sky emission lines; expressed in velocity units, this method is accurate to $\sim 1 \text{ km s}^{-1}$.

Pipeline software written in FORTRAN performs the reduction steps resulting in a wavelength- and flux-calibrated object spectrum along with an estimate of the noise, the sky transmission spectrum, and the 2-D unextracted slit image. The program also corrects for optical distortion, removes cosmic ray spikes and fixes bad pixels. The systematic accuracy of radiometric calibrations is $\sim 25\%$, ignoring telescope pointing uncertainties.

3. DESCRIPTION OF THE DATA

We observed IRS 9 in the lines of seven molecules at the settings summarized in Table 1. The selection of settings was motivated by the desire to include as many species and individual lines as possible given the severity of telluric atmospheric absorption at a particular setting. In particular, variations in the H_2O column over Mauna Kea often influenced the settings observed on a particular night, as many of the settings are sensitive to the intensities of telluric water lines. The systemic velocity of IRS 9 and motion of the Earth with respect to the Local Standard of Rest (LSR) were taken into account in choosing settings in which certain lines were shifted either into or away from nearby telluric absorption features, and in many cases required trade-offs in order to maximize efficiency and scientific return. Details of the observations for specific molecules follow.

3.1. C_2H_2

At Gemini, two lines in the acetylene ν_5 symmetric bending mode R branch were observed ($J=5,6$) along with approximately ten lines in the ν_5 Q branch. The R -branch lines are shown in Figure 1, while Figure 2 shows the Q -branch spectrum along with a simple LTE fit at a single rotational temperature of 100 K. This is discussed further in Section 4. Each acetylene line appears to consist of a single component. The $R(5)$ line is stronger than $R(6)$, as is expected due to its three-times-greater nuclear statistical weight. $R(6)$ appears to have a weak redshifted emission component at an LSR velocity of -54 km s^{-1} although this effect may be due to a poor fit to the continuum there. This component can be more easily seen in the top-center panel of Figure 10. Earlier IRTF observations of the ν_5 $R(1)$ line (wavenumber center 734 cm^{-1}) in 2001 June revealed a positive detection, but more importantly, higher J settings $R(15,16)$ (768 cm^{-1}) and $R(21,22)$ (780 cm^{-1}) were non-detections. The latter setting included the $^2P_{1/2} \rightarrow ^2P_{3/2}$ fine structure line of [Ne II]; consistent with the weak free-free emission observed toward IRS 9, this line was also a non-detection and is discussed in detail in Section 4.4. Finally, although the positions of lines of the $^{13}\text{C}^{12}\text{CH}_2$ isotopologue fell in some of our spectral settings, we did not detect any.

We also did not detect any lines in the $2\nu_5 - \nu_5$ “hot bands” of acetylene toward IRS 9. However, the hot-band lines in the spectral settings observed were not particularly favorable for detection, so we cannot draw any strong conclusions about the ν_5 population from these observations. The observation of $\nu_4 + \nu_5$ combination mode lines, shown in Figure 3, is of particular interest. This is because it allows a measurement of the C_2H_2 abundance with intrinsically weaker lines at different wavelengths from the ν_5 band lines. Since the $\nu_4 + \nu_5$ line strengths are ~ 5 times weaker than the ν_5 line strengths, absorption in the $\nu_4 + \nu_5$ lines may be dominated by high column density clumps of gas.

An energy level diagram indicating relevant vibrational levels of C_2H_2 is shown in Figure 4. The ν_5 mode is the symmetric bending mode in which the H atoms move together. In the ν_4 antisymmetric bending mode, they move in opposition, resulting in no dipole moment. Therefore, transitions that change the ν_5 quantum number are relatively strong, with $A_{ul} \sim 10 \text{ s}^{-1}$. These include the ν_5 , $2\nu_5 - \nu_5$ (between $2\nu_5$ and ν_5), and

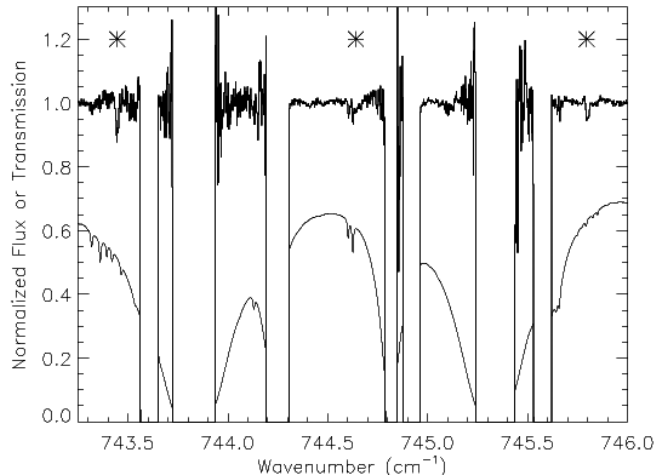


FIG. 1.— TEXES spectrum toward NGC 7538 IRS 9 showing features due to C_2H_2 and HCN. From left to right, the asterisks mark the positions of C_2H_2 ν_5 $R(5)$, HCN ν_2 $R(10)$, and C_2H_2 ν_5 $R(6)$. The upper trace shows the data, while the lower curve indicates the relative atmospheric transmission on the same scale. The HCN ν_2 $R(10)$ line is a detection despite the telluric lines at that wavenumber not completely dividing out. Gaps in the spectrum occur between echelon grating orders and where the telluric absorption is too great to be divided out.

$\nu_4 + \nu_5 - \nu_4$ (between $\nu_4 + \nu_5$ and ν_4) transitions. Transitions that only change ν_4 do not occur radiatively, but transitions changing both ν_4 and ν_5 , including $\nu_4 + \nu_5$ and $\nu_5 - \nu_4$, do occur although they are ~ 10 times weaker than those which only change ν_5 . Because of these selection rules, $7.6 \mu\text{m}$ photons can be absorbed in $\nu_4 + \nu_5$ lines, but molecules excited into the $\nu_4 + \nu_5$ level most often decay to the ν_4 level from which they can only decay to the ground vibrational level collisionally.

The ratio of the $\nu_4 + \nu_5$ $P(3)$ and $P(2)$ lines is difficult to determine due to poor telluric division, but appears to be less than the 9:2 ratio expected from their rotational and nuclear statistical weights. This probably indicates saturation of the lines in spite of their intrinsic weakness relative to the ν_5 transitions. We also note that the spectral region containing the $\nu_4 + \nu_5$ lines is contaminated by dense forests of HNO_3 and SO_2 lines, giving the superficial appearance of bad flat-fielding. Another difference between the ν_5 and the $\nu_4 + \nu_5$ transitions involves the shapes of the lines; the $\nu_4 + \nu_5$ lines are more asymmetric, with blue wings extending to nearly -90 km s^{-1} . In this regard they resemble the P Cygni-like shapes of the CO lines we observed toward IRS 9 (see Section 3.5). The $\nu_4 + \nu_5$ lines may result from absorption in a high column density component with a small covering factor that is strongly saturated. Or they may be formed in highly dust-obscured gas that is more visible at 7.6 than $13 \mu\text{m}$ due to the lower dust opacity off of the silicate absorption feature. The presence of the similar -80 km s^{-1} absorption feature in the CO lines supports this possibility. Alternately, the $\nu_4 + \nu_5$ line strengths relative to the ν_5 lines may be affected by other radiative transfer effects, as we discuss in Section 5.3.

3.2. HCN

Three of our spectral settings contain R -branch lines of the ν_2 bending mode of HCN with centers at 734 cm^{-1} ($6 \leq J \leq 7$), 745 cm^{-1} ($J=10$), and 761 cm^{-1}

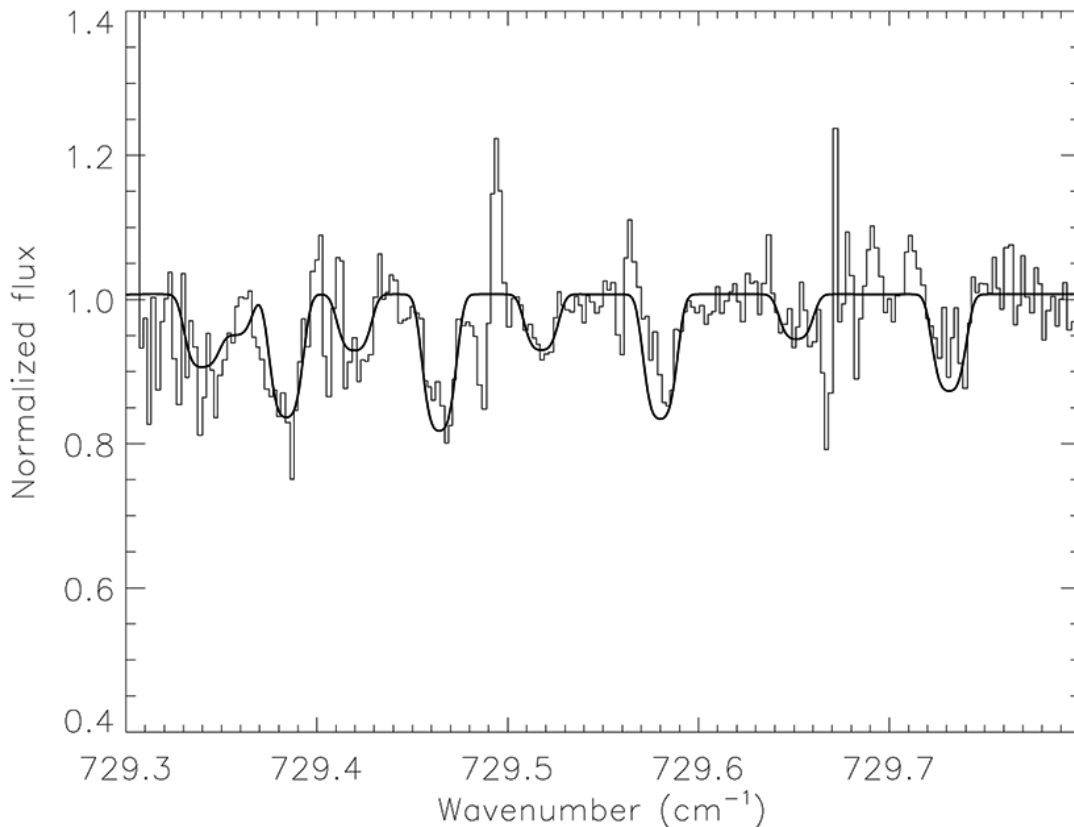


FIG. 2.— Spectrum of the ν_5 Q branch of C_2H_2 observed with TEXES toward NGC 7538 IRS 9 (thin lines). A fit is superimposed from an LTE pure absorption model with $T = 100$ K, $b = 1$ km s $^{-1}$, and $N(C_2H_2) = 2 \times 10^{15}$ cm $^{-2}$ (thick line). The synthetic spectrum has been convolved with the TEXES instrumental line shape function.

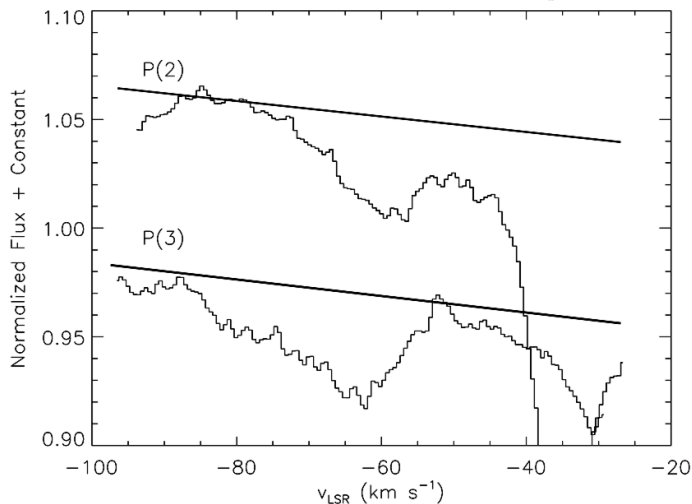


FIG. 3.— Spectra of two P -branch lines of the $\nu_4 + \nu_5$ combination mode of C_2H_2 shown on an LSR velocity scale. The two spectra have been offset vertically for clarity, and the continua are indicated with heavy straight lines. Both spectra are affected by strong, poorly-corrected telluric absorption at small blueshifts.

($J=16$). The $R(6)$ line fell into a region where atmospheric transmission is not favorable, and $R(7)$ appears to be a marginal detection because of a weaker, nearby atmospheric feature. $R(10)$ (Figure 1) is a probable detection. $R(16)$ falls close to the $R(13)$ line of C_2H_2 but is sufficiently separated from the acetylene line to prevent blending and is the only clean HCN line detected. The $R(15)$ line of HCN fell just outside the first order of the

761 cm $^{-1}$ setting. The S/N of the HCN observations is generally lower than that of C_2H_2 ν_5 R branch, but the HCN lines are clearly the broader of the two. The mean full width at half maximum (FWHM) of the four C_2H_2 R -branch lines is 6.5 ± 1 km s $^{-1}$, whereas the HCN $R(16)$ line FWHM is ~ 15 km s $^{-1}$. Finally, lines of the $H^{13}CN$ isotopologue are also found in these settings but not detected: $R(9)$ in 734 cm $^{-1}$ and $R(18)$ in 761 cm $^{-1}$. $R(18)$ falls into a gap between orders and $R(9)$ was affected by poor atmospheric transmission.

3.3. CH_4

We obtained data on three ν_4 R -branch lines of CH_4 : the $R(0)$ line at 1311.43 cm $^{-1}$ and the two $R(2)$ lines at 1322.083 cm $^{-1}$ and 1322.152 cm $^{-1}$. In addition, the settings for these lines contained the positions of the $R(1)$ and $R(4)$ lines of $^{13}CH_4$, but both of these lines were clear non-detections. The $^{12}CH_4$ lines show clear asymmetries indicative of multiple velocity components. Each feature has a principal component corresponding to greatest absorption at $v_{LSR} = -65.0$ km s $^{-1}$ with additional components on the red wing of this feature separated in velocity by up to 6.5 km s $^{-1}$. The spectra are plotted in Figure 5.

3.4. NH_3

Two spectral settings observed contain a total of ten P -branch lines of ammonia, three in the ν_2 “umbrella” mode antisymmetric $P(4,K)$ group and seven in the ν_2 symmetric $P(7,K)$ group. Of the possible lines in these

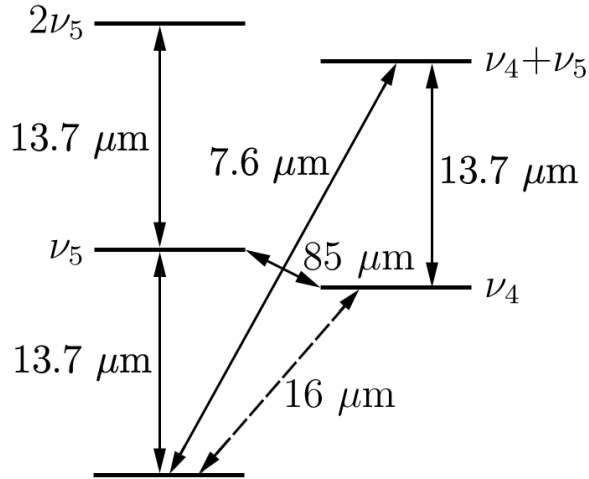


FIG. 4.— Energy level diagram for some mid-infrared-active, ro-vibrational transitions of C_2H_2 . Radiatively-allowed transitions are denoted with solid lines and collisionally-allowed-only transitions with dashed lines. The excited vibrational levels are split by interactions between the rotational angular momentum and the angular momentum of the bending-mode vibration; however, the energy differences between these levels are too small to be seen at the resolution of the diagram.

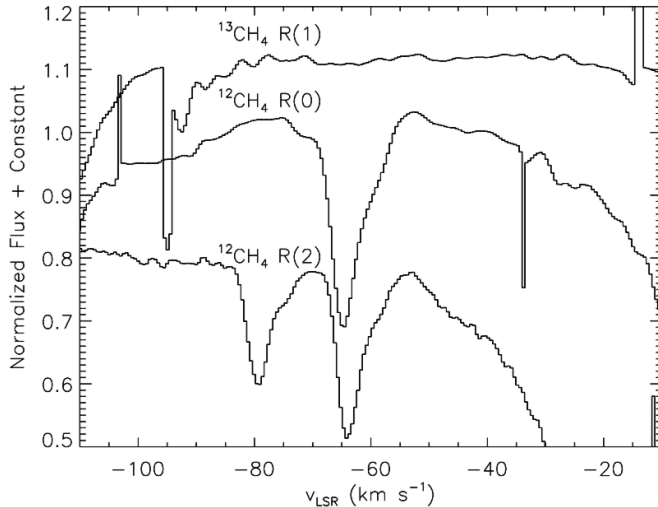


FIG. 5.— Spectra of three features in the ν_4 band of CH_4 toward IRS 9, shown on an LSR velocity scale. For the spectrum containing the $R(2)$ lines, the velocity scale has been set to the LSR value for the stronger component (rest frequency 1322.085 cm^{-1}). The individual spectra have been offset vertically for clarity.

settings, only $aP(4,2)$ was not observed because of its proximity to a strong telluric atmospheric feature. Almost all of the NH_3 lines share a characteristic P Cygni-like profile, with an emission peak near the systemic velocity, around -60 km s^{-1} , and weak blueshifted absorption separated in velocity by 2.5 to 10 km s^{-1} . This may be compared to the -56.5 and -59.8 km s^{-1} absorption components seen in mid-infrared NH_3 absorption transi-

tions toward IRS 1 (Knez et al. 2009) and a value of -60.0 km s^{-1} for the $(3,3)$ rotational transition of NH_3 in Very Large Array (VLA) observations of IRS 1 at centimeter wavelengths (Henkel et al. 1984). The degree of absorption toward IRS 9 relative to emission varies as well; the highest K lines in the $sP(7,K)$ series show no absorption at all despite the fair S/N of the data. Finally, we observed a setting near 930 cm^{-1} containing the positions of the ν_2 aQ-branch lines ($1 \leq J \leq 9$); the data show weak absorption features at some of the indicated points in Figure 6. For comparison, a similar plot of the $sP(7,K)$ branch spectrum is also shown in the figure. We did not observe lines in the ν_2 R-branch of NH_3 because of their proximity to regions of significant telluric atmospheric absorption. Also, the R-branch lines are located in a spectral region of rapidly increasing absorption due to the $9.7 \mu\text{m}$ silicate dust feature. This fact significantly impacts our interpretation of the NH_3 P- and R-branch lines, as we discuss in Section 5.3.

3.5. CO

We observed two settings in the $5 \mu\text{m}$ $v = 1-0$ band of CO, which included a number of P-branch lines of several isotopologues. These settings were centered approximately at 2055 cm^{-1} and 2085 cm^{-1} . Strongest among the detected features were the ^{12}CO lines with $J = 14, 15, 21$ and 22 . Each line exhibited a strong P Cygni profile with emission and absorption components separated by $\sim 20 \text{ km s}^{-1}$. Each component was broader than those of any line of any other molecular species we observed, with line widths of $10\text{--}20 \text{ km s}^{-1}$. In each case the strength of the emission was approximately equal to the strength of the absorption. The features are consistent with an expansion or outflow and have been previously interpreted as such by Mitchell & Hasegawa (1991). Observations of IRS 1 were also made at these settings; we find a different profile altogether with a double-bottomed absorption feature in ^{12}CO , weaker absorption on the blue wing and a slight emission bump on the red wing of each feature. Spectra of both objects are shown in Figure 7.

Four ^{13}CO P-branch lines with $J = 3, 4, 11$, and 12 were also detected in the spectrum of IRS 9. Of these, $P(11)$ was a clear detection near the middle of a spectral order and away from contamination by telluric atmospheric lines, $P(12)$ was noisy and found very near the edge of a spectral order, and $P(8)$ and $P(9)$ are clearly detected, but blended slightly with two nearby C^{18}O features. Additional ^{13}CO lines, $P(11)$ and $P(12)$, are located in the spectral setting at 2055 cm^{-1} but occur in regions of strong telluric atmospheric absorption and thus were not detected. We also searched for lines attributable to C^{17}O but did not confidently detect any; this is not surprising given the low expected abundance of C^{17}O . For reference, Smith et al. (2009) quote $N(^{12}\text{C}^{16}\text{O})/N(^{12}\text{C}^{17}\text{O})$ of 2800 ± 300 toward the T Tauri star VV CrA, obtained at a spectral resolution similar to our TEXES observations. They compare this to the Local ISM value of 2005 ± 155 (Wilson 1999).

The CO lines seen toward both IRS 9 and IRS 1 have different line shapes in ^{12}CO and ^{13}CO . In the case of IRS 1, the ^{13}CO features show only single components in the deepest absorption, while toward IRS 9 the ^{13}CO

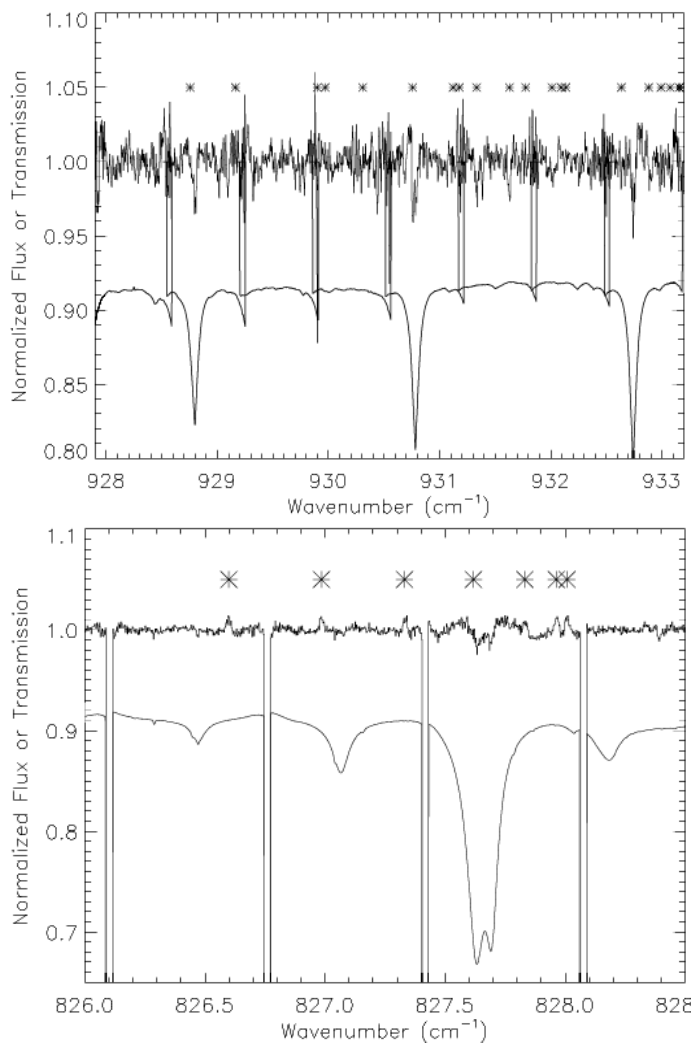


FIG. 6.— TEXES spectra toward NGC 7538 IRS 9 at the position of the $^{14}\text{NH}_3$ ν_2 antisymmetric Q branch (top panel) and symmetric $P(7,K)$ branch (bottom panel). The upper curve in each plot shows the data, while the lower curve in each shows the relative atmospheric transmission on the same scale. Asterisks mark positions of the NH_3 lines. Both spectra have been corrected for the Earth’s motion relative to the LSR, and the position of the NH_3 line markers shifted to the systemic LSR velocity of approximately -60 km s^{-1} .

$P(11)$ line and ^{12}CO lines show the pronounced P Cygni profiles. The ^{13}CO lines in the 2085 cm^{-1} setting clearly show multiple components suggestive of P Cygni profiles in which emission is weak relative to absorption, and there are at least two distinct absorption units. In turn, the strongest absorption components show a weak, double-bottomed structure like the ^{12}CO lines in IRS 1. The multiple absorption units may appear in the profiles of the ^{12}CO features in the same spectral setting; weak “shoulders” on the blue wing of the absorption component are consistent with additional absorption suffering contamination by the primary components. Such shoulders are also noticeable on the red wings of the emission features of ^{12}CO . These details can be seen in Figure 8.

The detection of ^{13}CO toward IRS 9 with line depths comparable to those of ^{12}CO strongly indicates saturation of both isotopologues. Useful columns may however be obtained from rarer variants such as C^{18}O , but

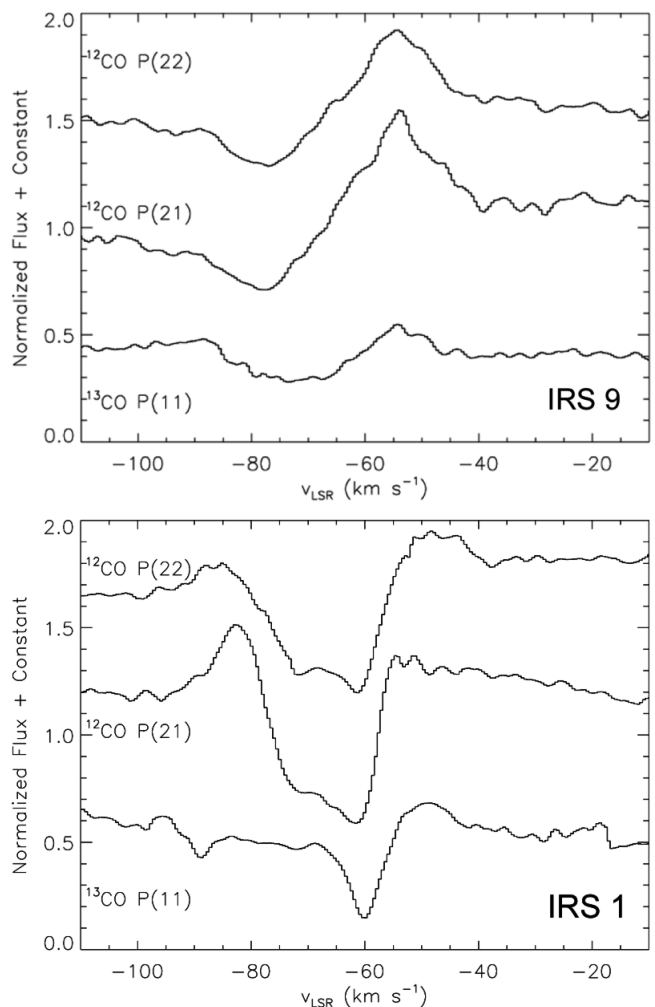


FIG. 7.— The upper panel shows spectra of two P -branch lines of ^{12}CO and one P -branch line of ^{13}CO toward IRS 9 obtained with TEXES. The lower panel shows the same spectral settings observed toward IRS 1. In both cases the motion of the Earth with respect to the LSR has been removed from the velocity scales.

these lines are too weak to be measured accurately in the present spectra. It would also be of interest measure the $^{12}\text{C}/^{13}\text{C}$ ratio toward IRS 9 for comparison with the value of 45 toward IRS 1 reported by Knez et al. (2009). As we did not observe the ^{13}C isotopologue of C_2H_2 in our spectra of IRS 9, the CO observations would offer a way to measure the ratio if the lines were not saturated.

3.6. OCS

The ^{12}CO setting at 2055 cm^{-1} we observed also contains the positions of a number of P -branch lines of the ν_2 stretching mode of OCS ($12 \leq J \leq 22$), but we cannot confidently claim detections of any of them toward IRS 9. Periodic structure in the continuum of our $5 \mu\text{m}$ spectra may be broad OCS features, although it is more likely the structure results from flat-fielding residuals. For reference, we have overplotted the locations of the rest-frame OCS line centers over the data in Figure 9. It is unclear whether we should have observed lines of OCS in the gas phase, given the strong detections of CO nearby. Low temperatures may inhibit desorption of OCS from grains, but the theoretical grounds for predicting OCS abundances are questionable (Doty et al.

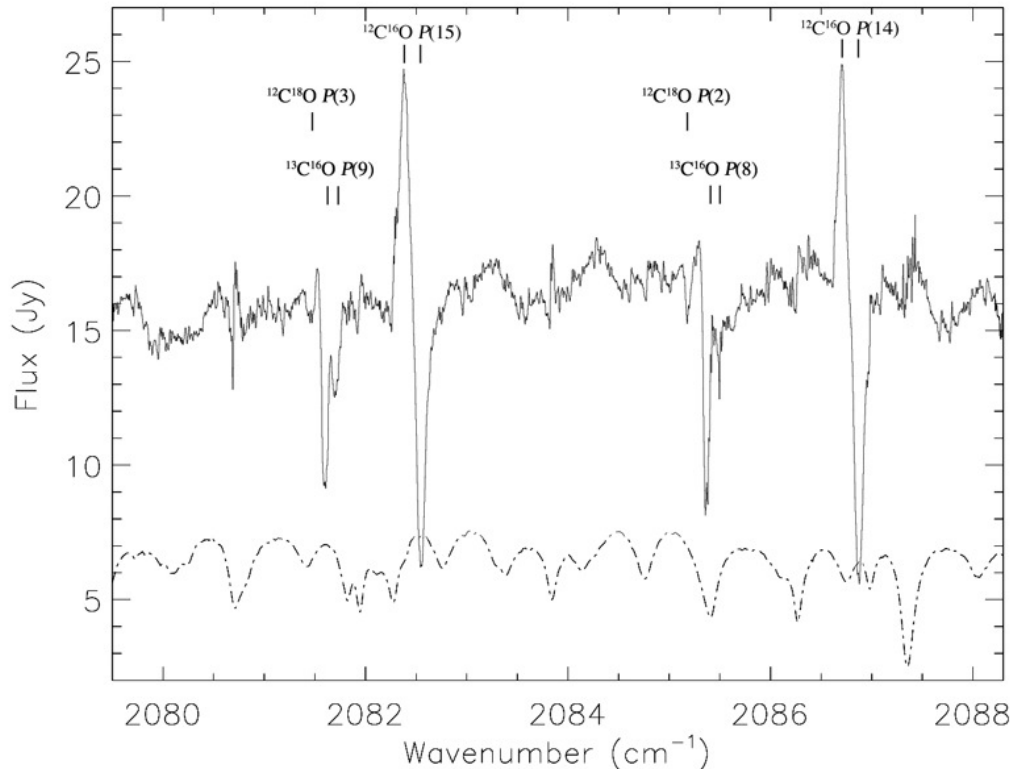


FIG. 8.— TEXES spectrum toward NGC 7538 IRS 9 centered near the $5\ \mu\text{m}$ rotational lines of several isotopologues of CO. The data are shown in the upper trace (solid line) and the relative atmospheric transmission in the lower trace (dot-dashed line) multiplied by a factor of 10. Some line identifications are shown.

2004). The signature of an outflow in the spectra of CO hints at possible excitation through shocks. We might expect that OCS would be similarly excited but there is no support for this in our data. This is contrasted with the results of Evans et al. (1991), who detected gas-phase OCS toward Orion IRC2; for reasons given in Section 5, it is impossible to know whether this is due to temperature differences or real abundance variations between the IRS 9 and IRS 1.

3.7. HNC O

We looked for ν_4 bending mode HNC O IL-0 P -branch lines toward IRS 9 in the TEXES spectral setting that included $s(7, K)$ lines in the P branch of ammonia. None of the three lines included in this setting ($2 \leq J \leq 4$) was detected, although there may be HNC O Q -branch absorption at $827.7\ \text{cm}^{-1}$, obscured by telluric interference. Given the strong observed ice absorption at $4.62\ \mu\text{m}$ toward IRS 9 (Pendleton et al. 1999), now commonly attributed to OCN^- , and a lack of gas-phase HNC O absorption, it seems most likely that temperatures in the outer envelope of IRS 9 are sufficiently low to prevent OCN^- desorbing from grains in detectable quantities. Alternately, gas-phase OCN^- might exist at H (or H_2) abundances insufficient to result in considerable HNC O production.

4. ANALYSIS AND MODELING

We fit synthetic spectra to the TEXES data to derive excitation parameters and column densities for the various molecular species observed. A two-step process

was devised in which rough estimates of the parameters were made from “by eye” fits of synthetic spectra to our data; these values were then used as the initial guesses in our fitting program which iterates values of hundreds of parameters to achieve the best fit to the data by minimizing χ^2 . While our χ^2 code is robust in finding solutions even when starting far from the global minimum, convergence is achieved faster by using good initial guesses. This approach was best suited to our method of spectral analysis.

4.1. Simple Synthetic Spectrum Fitting

In order to get approximate parameters for the gas-phase lines we observed, we first fit synthetic spectra to the data using a fitting method commonly referred to as “chi by eye.” This approach allowed us to get quick initial estimates of parameters such as excitation temperature and column density for the various molecules that were used as inputs to a more sophisticated non-linear least squares code to obtain optimal values of the parameters. The simple method also allowed us to determine the number, width and velocities of line components.

The equivalent width of an optically-thin line is the product of the line strength, α , and the column density, N . We obtained values of α for each molecular species observed from the GEISA database¹ of laboratory molecular data (Jacquinet-Husson et al. 2005). The GEISA coefficients are specified for a fiducial temperature $T_o = 296\ \text{K}$; we corrected them to an assumed temperature T

¹ http://ether.ipsl.jussieu.fr/ether/pubipsl/geisa_iasi_raie_frame_2003_uk.jsp

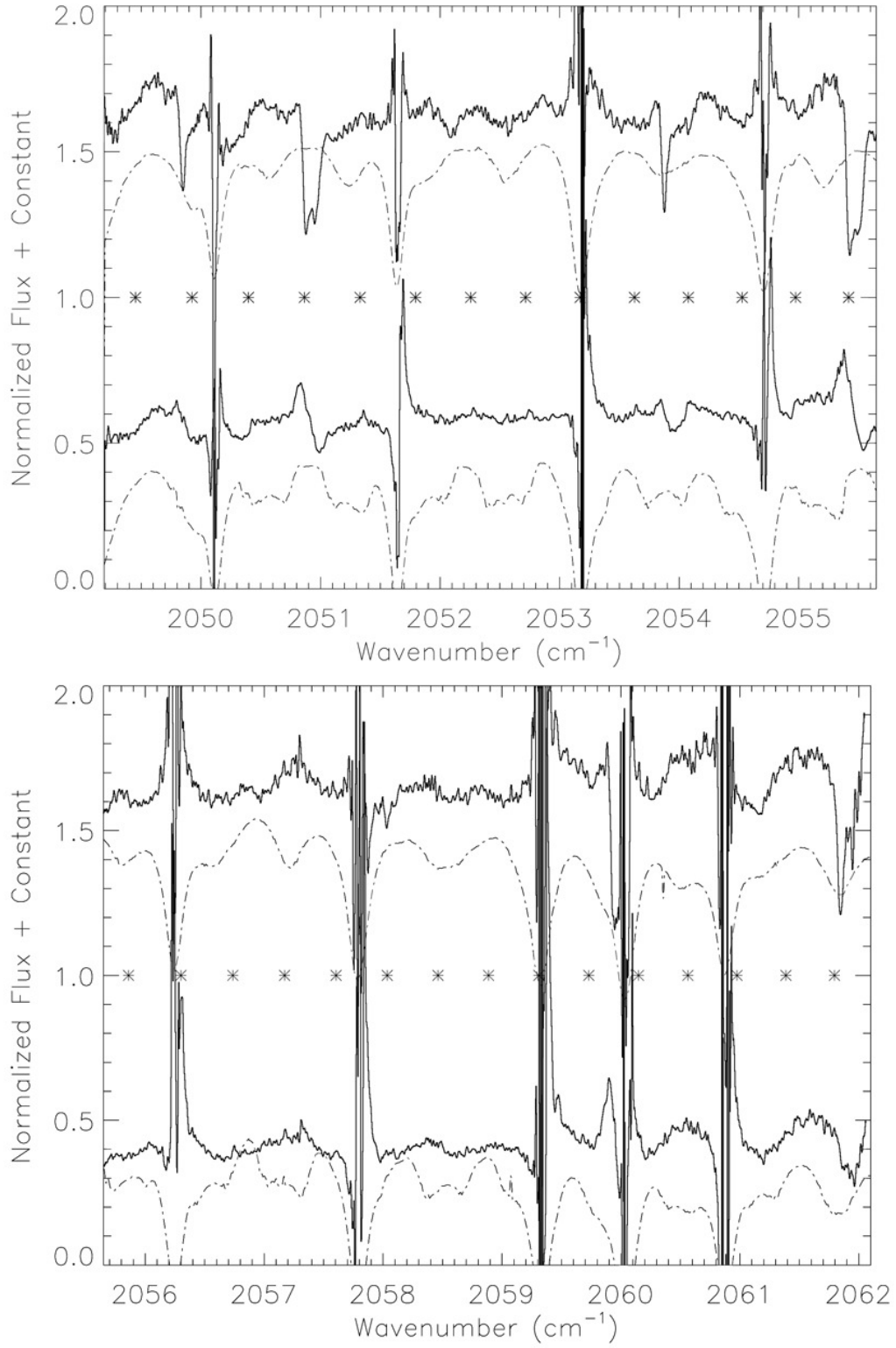


FIG. 9.— The TEXES spectra of NGC 7538 IRS 1 (top trace in each panel) and IRS 9 (bottom trace) near features in the P -branches of ^{12}CO and ^{13}CO ; heavy solid lines are the object spectra and light dot-dash lines are the sky spectra offset for clarity. The wavenumber scale has been corrected to a velocity reference at rest in each object's frame, and the line centers of OCS features, marked with asterisks, have been velocity shifted appropriately.

via the formula

$$\alpha(T) = \alpha(T_o) \exp \left[E_l \left(\frac{1}{k_B T_o} - \frac{1}{k_B T} \right) \right] \left(\frac{T_o}{T} \right)^\beta \quad (2)$$

where E_l is the energy of the lower state of a transition, k_B is Boltzmann's constant, and $\beta = 1$ for linear molecules and $\frac{3}{2}$ for non-linear molecules. Here the temperature dependence of the vibrational partition function and the stimulated emission correction are neglected as both are of order unity. For Gaussian lines, the line-center optical depth is calculated according to

$$\tau_l = \frac{N\alpha}{b\sqrt{\pi}} \quad (3)$$

where b is the Doppler b parameter in wavenumber units. The run of optical depth with frequency over the line is computed with

$$\tau(w) = \tau_l \exp \frac{-(w - w_o)^2}{b^2} \quad (4)$$

in which w is the wavenumber and w_o is the line-center wavenumber. The optical depths are converted to intensities by taking the exponential of $\tau(w)$; the continua are normalized to unity. We assume that the rotational states are in local thermodynamic equilibrium (LTE). Our radiative transfer model also assumes pure absorption for fitting the observed absorption lines. Re-emission by the absorbing molecules is neglected. Emission features are treated as separate components, which are modeled as negative absorption. Populations of excited vibrational states are not calculated. Temperatures are arbitrarily held fixed at 200 K; this value is consistent with temperatures around IRS 1 derived from TEXES data. The exception was the Q -branch of C_2H_2 where a number of observed J states allowed the temperature to be constrained to 100 K. Column densities and velocities derived using this method are shown in Table 2. Fits to the individual lines are shown in Figures 10 through 12.

4.2. χ^2 Minimization Spectral Fitting

Our spectral fitting code, *FITSPEC*, uses the Levenberg-Marquardt method of inverting the second derivative matrix to minimize χ^2 as in Bevington & Robinson (2003) for up to four unique absorbing or emitting units, or “clouds,” along the line of sight. The 1-D radiative transfer model assumes rotational LTE of all molecules with a single excitation temperature. As in our “by eye” fitting, re-emission by absorbing molecules is neglected and emission features are modeled as negative absorption. The total number of free fitting parameters depends on the number of molecules in the fit and the number of echelon orders in a particular spectral setting. Although there is a large number of free parameters in the fit, most of them are “uninteresting,” such as baseline coefficients. The “interesting” parameters in the fit are Doppler shift, Doppler b parameter, column density, temperature, and covering factor for each cloud and for each molecule. The details of the program are discussed in Knez et al. (2009). For our observed molecules, we chose a number of absorbing clouds consistent with the number of velocity components seen in the TEXES spectra, allowing *FITSPEC* to

independently determine their parameters. When available, we included data indicating the non-detection of certain lines, such as high- J C_2H_2 near $13 \mu m$, to better constrain the fit parameters.

The success of *FITSPEC* requires accurate molecular transition data to compute synthetic spectra against which to fit the data. We relied on the calculations in Evans et al. (1991), giving information including the line center, the energy of the transition either in units of $J(J+1)$ (which is then multiplied by the molecule's rotational constant B) or E in units of cm^{-1} , and the rotational line strengths (Hönl-London factors) in units of $cm^{-1} (10^{16} cm^{-2})^{-1}$. The program uses guesses of the starting parameter values to compute the initial model spectrum, and then iterates the fit by varying those values until a convergence criterion is satisfied. The requirement is

$$\sum_i \left(\frac{\Delta p_i}{\sigma_{p_i}} \right)^2 \leq 0.01 \quad (5)$$

where Δp_i is the change in the value of parameter i and σ_{p_i} is the uncertainty in the parameter i , allowing all other parameters to vary.

The output of the program consists of the best-fit parameter values, uncertainty estimates and a synthetic spectrum having the properties of the best fit. The intensity of the synthetic spectrum is calculated by taking $I = I_o e^{-\tau}$ of each absorbing unit along the line of sight and adding any optically thin emission as

$$I_{obs} = I_0 \prod_{n=1}^N [1 - C_n(1 - e^{-\tau_n(\nu)})] + S_{em} C_{em} \tau_{em}(\nu) \quad (6)$$

where N is the number of absorbing clouds in the fit, I_0 is the intensity of radiation from the background continuum source (set to unity since the spectra are normalized), $C_n = \Omega_n/\Omega_0$ and τ_n are the covering factors and optical depths of the absorbing clouds, S_{em} and $C_{em} = \Omega_{em}/\Omega_0$ are the source function and covering factor of the emitting material, and Ω_0 is the solid angle of the continuum source. The right-hand term represents an optically-thin emitting cloud; since it only depends on the product of S_{em} and C_{em} , we scale C_{em} by S_{em}/I_0 and drop S_{em} from the formula. The product term in Equation 6 assumes that the absorbing clouds overlap randomly. That is, if two clouds each have covering factors of 0.1, 1/10 of each lies along the line of sight to the other.

Knez et al. (2009) used the same approach and note this results in greater absorption in saturated lines than would result if the optical depths were added first and then the transmission spectrum were calculated from the optical depth spectrum. The resulting fits to the data are better than those obtained by summing the optical depths first and then taking the exponential. This fitting approach, however, is not without drawbacks. Emission lines may not be thermally excited, but rather could be due to resonantly scattered light; we have no way of discriminating between these processes. We assume that upward transitions in the molecules we observe are not followed by re-emission of that radiation; in reality, most absorptions should be followed by re-emission, although in a direction uncorrelated with the incoming radiation,

TABLE 2
MOLECULAR ABSORPTION PARAMETERS DERIVED FROM SIMPLE SYNTHETIC SPECTRUM FITTING IN LTE

Molecule/Band	Line	n ^a	T (K)	v_1^b (km s ⁻¹)	v_2	v_3	b_1 (km s ⁻¹)	b_2	b_3	N_1^c (cm ⁻²)	N_2	N_3
C ₂ H ₂ ν_5	R(1)	1	100	-62.0	—	—	3.3	—	—	1.0(15)	—	—
	R(5)	1	100	-62.0	—	—	2.4	—	—	4.5(14)	—	—
	R(6)	1	100	-65.5	—	—	2.4	—	—	4.5(14)	—	—
	Q branch	1	100	—	—	—	1.0	—	—	2.0(15)	—	—
C ₂ H ₂ $\nu_4 + \nu_5$	P(2)	1	100	-61.0	—	—	6.8	—	—	2.5(16)	—	—
	P(3)	1	100	-62.0	—	—	5.7	—	—	1.0(16)	—	—
HCN ν_2	R(16)	1	200	-64.5	—	—	7.9	—	—	1.2(16)	—	—
¹² CH ₄ ν_4	R(0) ^d	4	200	-65.0	-54.0	-61.0	2.3	2.7	4.6	7.5(16)	-4.0(15)	6.0(16)
	R(2)	2	200	-65.5	-61.5	—	2.3	2.5	—	6.0(16)	3.3(16)	—
	R(2)	—	—	—	—	—	2.5	2.3	—	6.3(16)	1.5(16)	—
¹⁴ NH ₃ ν_2	aP(4,0)	2	200	-64.0	-70.0	—	4.6	3.9	—	-2.5(15)	5.0(14)	—
	aP(4,1)	2	200	-62.0	-72.0	—	5.3	3.9	—	-3.8(15)	2.5(14)	—
	aP(4,3)	2	200	-62.0	-70.0	—	2.8	3.9	—	-1.0(15)	1.0(15)	—
	sP(7,0)	2	200	-60.0	-69.0	—	2.9	2.9	—	-1.5(15)	1.2(15)	—
	sP(7,1)	2	200	-62.0	-70.0	—	1.8	4.0	—	-2.0(15)	7.5(14)	—
	sP(7,2)	2	200	-61.0	-62.5	—	3.3	2.9	—	-6.0(15)	4.5(15)	—
	sP(7,3)	2	200	-58.5	-64.0	—	1.8	2.9	—	-1.2(15)	1.5(15)	—
	sP(7,4)	2	200	-60.5	-63.0	—	1.8	1.8	—	-2.5(15)	2.0(15)	—
	sP(7,5)	1	200	-56.0	—	—	4.4	—	—	-7.0(15)	—	—
	sP(7,6)	1	200	-57.0	—	—	4.4	—	—	-4.0(15)	—	—
¹² C ¹⁶ O v=1-0	P(14)	3	200	-55.0	-64.5	-88.5	2.4	6.5	7.2	-1.0(16)	-1.1(17)	3.0(17)
	P(15)	3	200	-51.5	-63.5	-88.5	3.6	9.4	7.2	-2.5(16)	-2.2(17)	3.0(17)
	P(21)	2	200	-54.5	-77.5	—	10.9	10.9	—	-8.0(17)	5.5(17)	—
	P(22)	2	200	-55.0	-77.0	—	8.0	10.2	—	-6.0(17)	6.0(17)	—
¹³ C ¹⁶ O v=1-0	P(3)	3	200	-60.5	-70.5	-89.0	4.3	4.3	1.4	-9.0(17)	8.0(18)	7.0(17)
	P(4)	3	200	-61.5	-71.5	-86.0	3.6	5.0	6.5	-7.0(17)	6.8(18)	3.4(18)
	P(11)	2	200	-54.0	-75.0	—	5.8	10.2	—	-8.0(17)	2.7(18)	—
	P(12)	2	200	-54.0	-73.0	—	5.9	10.2	—	-3.5(17)	6.5(18)	—
¹² C ¹⁸ O v=1-0	P(2)	2	200	-68.0	-72.0	—	2.9	3.6	—	7.5(18)	7.5(18)	—
	P(3)	2	200	-66.0	-71.0	—	4.3	4.3	—	3.0(18)	3.0(18)	—

^a Number of velocity components fit.

^b All velocities are referred to the LSR.

^c Column densities are given in the form $A(B) = A \times 10^B$. Negative values of A indicate components in emission. “Column density” includes all rotational states, assuming $T = 200$ K.

^d The best fit to this line required four components. The remaining component not in the table is $v_4 = -71.5$ km s⁻¹, $b_4 = 2.3$ km s⁻¹, $N_4 = 5.0(15)$ cm⁻²

so the extent to which re-emission fills in the absorption depends on the distribution of the re-emitting gas around the continuum source.

The results of our FITSPEC runs on the IRS 9 data are summarized in Table 3 where they are compared with results from fitting TEXES data of IRS 1 with the same procedure in Knez et al. (2009).

4.3. Column Density Upper Limits For Non-Detections

We calculate column density upper limits from the lines that were not detected in the spectrum of IRS 9 by convolving the data with a Gaussian line shape function whose properties are consistent with those of detected lines. We begin by assuming a normalized spectrum containing a single absorption line has the functional form

$$F_\nu = 1 - ae^{-x^2/b^2} \quad (7)$$

where $a > 0$ is the depth of the line and b is the Doppler b parameter. The observed spectrum is convolved with a Gaussian of the form ne^{-x^2/b^2} where n is a normalization factor here chosen to be $1/\sqrt{\pi}b$ such that the integral of the Gaussian is unity. The convolution is given by

$$F_\nu \otimes ne^{-x^2/b^2} = \int F_\nu ne^{-x^2/b^2} d\nu, \quad (8)$$

where $x = \nu - \nu'$. The upper limit of the equivalent width, W_{max} , of a line not detected is proportional to the peak value of the convolved spectrum according to

$$W_{max} = \sqrt{\pi}Pb \quad (9)$$

in which P is the peak. We used $b = 3.5$ km s⁻¹ in making the calculations, a value consistent with those derived by spectral fitting methods described previously. The column density limit for species X was then computed from

$$N(X) = \frac{W_{max}}{\alpha} \quad (10)$$

where α is the line strength. We used the GEISA line strengths where available and corrected them to a uniform temperature of 200 K as described previously. Results for our non-detections are given in Table 4.

We derived upper limits for all lines of a particular molecular species available in a given spectral setting, as indicated in the table by the number of lines parameter n . In cases where $n > 1$, the equivalent width and column density upper limits given are simple averages. All limits are given with respect to the total number of molecules along the column, irrespective of isotopic composition except for ¹³CH₄; here we have divided the derived column density by the terrestrial ¹²C/¹³C = 89 ratio built into GEISA to recover the column density specifically for

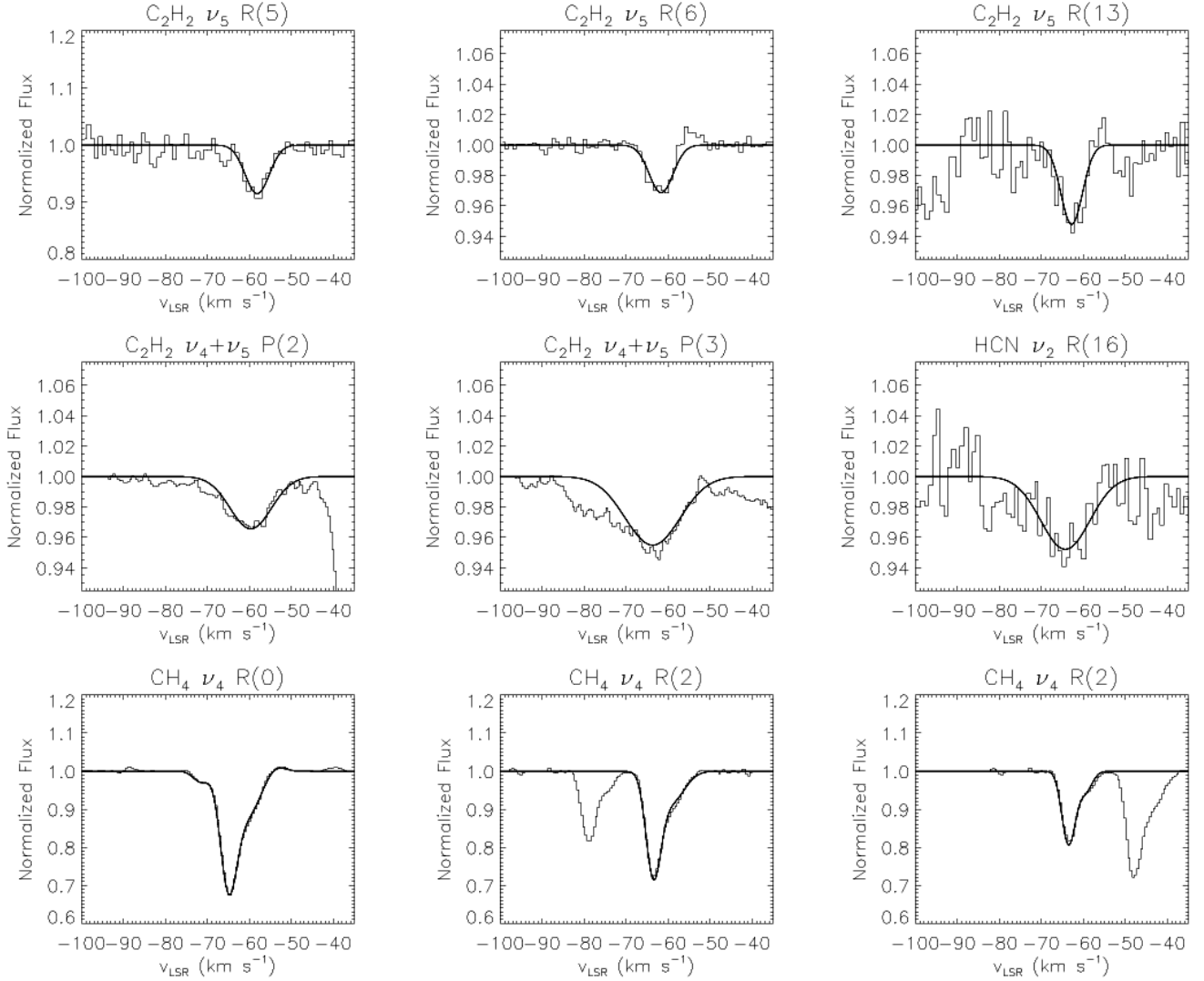


FIG. 10.— TEXES spectra of NGC 7538 IRS 9 (thin lines) and synthetic spectra computed using the simple LTE fitting method described in Section 4.1 (thick lines). The continuum in each spectrum has been normalized to unity and the spectra are plotted on a velocity scale referred to the LSR. Features shown are C_2H_2 ν_5 $R(5)$, $R(6)$ and $R(13)$; C_2H_2 $\nu_4 + \nu_5$ $P(2)$ and $P(3)$; HCN ν_2 $R(16)$; and CH_4 ν_4 $R(0)$ and the two components of $R(2)$.

$^{13}\text{CH}_4$.

4.4. An Upper Limit to the Ionizing Flux in IRS 9

The $^2P_{1/2} \rightarrow ^2P_{3/2}$ fine structure line of $[\text{Ne II}]$ at 780.42 cm^{-1} is often used as a diagnostic of ionized gas at our wavelengths. The non-detection of this line in the spectrum of IRS 9 places a constraint on the amount of ionizing flux being emitted by the embedded central source.

The Lyman continuum luminosity (N_{Lyc}) can be calculated from the $[\text{Ne II}]$ luminosity since both are proportional to the square of the density integrated over volume. The $[\text{Ne II}]$ volume emissivity is given by

$$j_{[\text{Ne II}]} = n_e n_{\text{Ne}^+} q_{lu} h\nu \frac{n_c}{n_e + n_c} \quad (11)$$

where n_e is the electron density, n_{Ne^+} is the density of ionized neon, q_{lu} is the rate coefficient for collisional excitation ($q_{lu} h\nu = 8.4 \times 10^{-22} \text{ erg s}^{-1} \text{ cm}^3$), and $n_c = A_{ul}/q_{ul}$ ($= 4.9 \times 10^5 \text{ cm}^{-3}$ at an electron temperature

$T_e = 10^4 \text{ K}$) is the critical density in the two-level approximation, using the values in Osterbrock & Ferland (2006). The hydrogen recombination rate per unit volume is given by $n_{\text{recomb}} = \alpha_B n_e n_p$ where α_B ($= 2.6 \times 10^{-13} \text{ cm}^3 \text{ s}^{-1}$) is the recombination rate for “case B” recombination, and n_p is the proton density. The ratio of N_{Lyc} to the $[\text{Ne II}]$ luminosity is then given by

$$\frac{N_{\text{Lyc}}}{L_{[\text{Ne II}]}} = \frac{n_{\text{recomb}}}{j_{[\text{Ne II}]}} = \frac{n_{\text{H}^+}}{n_{\text{Ne}^+}} \frac{\alpha_B}{q_{lu} h\nu} \quad (12)$$

assuming $n_e \ll n_c$. The Ne^+/H^+ ratio is given by the product of the Ne abundance and the singly-ionized fraction for Ne, f_+ . We take the value of $f_+ = 0.3$ from Figure 3 in Ho & Keto (2007) using for the radiation temperature that of a ZAMS star with IRS 9’s luminosity ($T_r = 26200 \text{ K}$) and a ratio of the dilution factor for blackbody radiation, Γ , to the electron density of $\Gamma/n_e = 10^{-19} \text{ cm}^3$. From this we adopt a Ne^+/H^+ ratio to be 3×10^{-5} .

Without knowing the width of the $[\text{Ne II}]$ line, it is

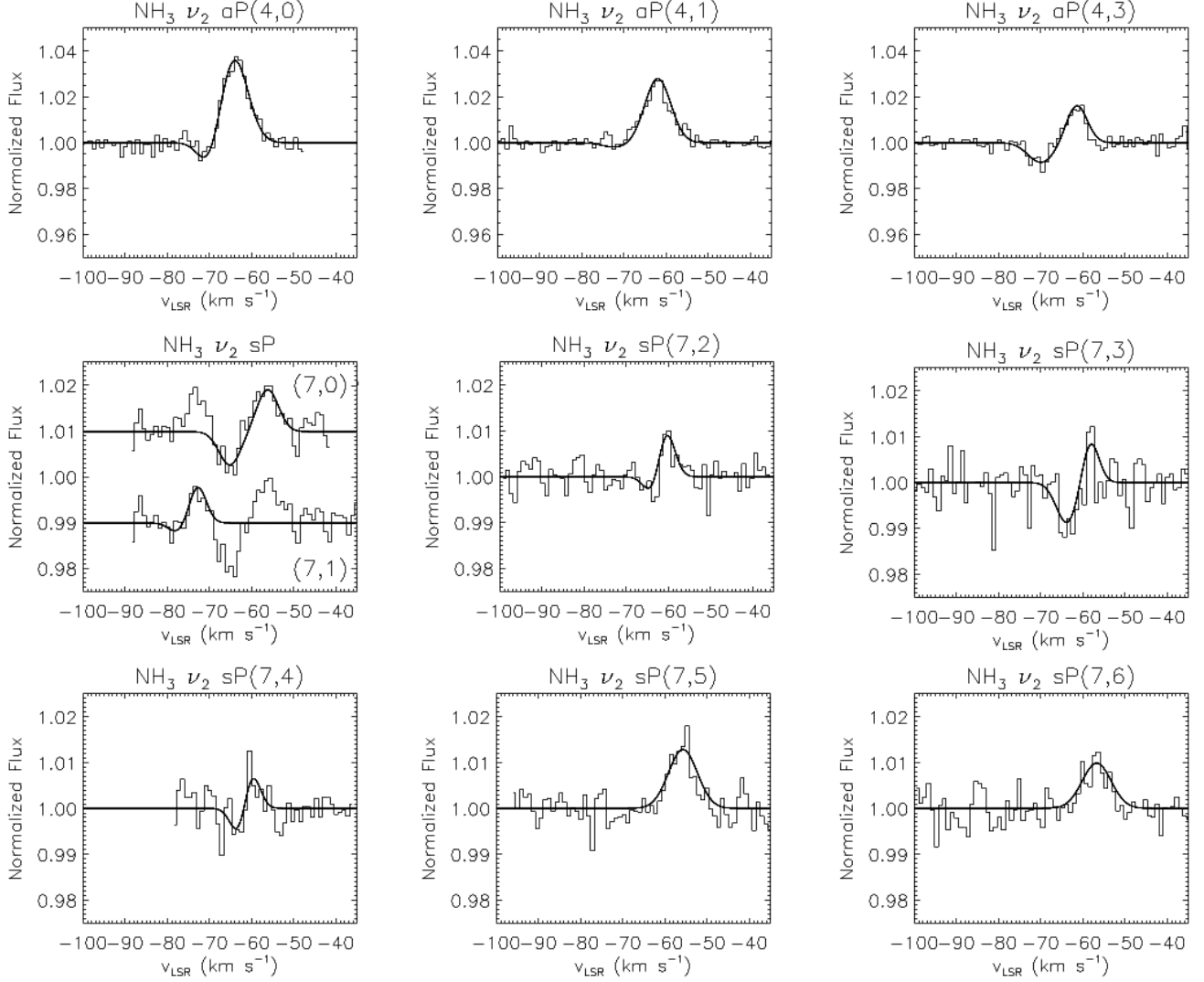


FIG. 11.— TEXES spectra of NGC 7538 IRS 9 (continued). Features shown are $\text{NH}_3 \nu_2$ aP(4,0), aP(4,1), aP(4,3), sP(7,0), sP(7,1), sP(7,2), sP(7,3), sP(7,4), sP(7,5) and sP(7,6). The sP(7,0) and sP(7,1) lines are shown together in a single panel using the velocity scale for sP(7,0), and are offset slightly for clarity.

difficult to place a strong upper limit on its flux, but an emission equivalent width of $W_\nu = 0.01 \text{ cm}^{-1}$ can be ruled out unless the line is wider than 0.2 cm^{-1} , or 77 km s^{-1} . This corresponds to a line flux limit of $F_{[\text{NeII}]} < 2.0 \times 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2}$ and a luminosity limit of $L_{[\text{NeII}]} = 4\pi d^2 F_{[\text{NeII}]} < 1.7 \times 10^{32} \text{ erg s}^{-1}$. We then calculate $N_{\text{Ly}\alpha} < [3.3 \times 10^4 \times 2.6 \times 10^{-13} / 8.4 \times 10^{-22}] \times 1.7 \times 10^{32} = 1.7 \times 10^{45} \text{ photons s}^{-1}$. For reference, Tielens (2005) gives the Lyman continuum luminosity of a main sequence B0 star as $N_{\text{Ly}\alpha} = 1.4 \times 10^{48} \text{ photons s}^{-1}$. Apparently IRS 9 emits a small fraction of the Lyman continuum luminosity of a main-sequence star of its bolometric luminosity. The radiation environment of IRS 1 is considerably more intense, by comparison. From infrared and millimeter data, the exciting source is inferred to be of spectral type earlier than O7.5, total luminosity $L > 8 \times 10^4 L_\odot$, and Lyman continuum luminosity $N_{\text{Ly}\alpha} > 10^{48} \text{ s}^{-1}$ (Hackwell et al. 1982; Akabane et al. 2001; Lugo et al. 2004). The inferred radiation field of IRS 1 is more intense than that of IRS 9 and results in a different set of conditions under which chemistry proceeds.

5. INTERPRETATION OF THE SPECTROSCOPIC RESULTS

5.1. Assumptions and Caveats

Interpretation of our results requires caution, given the underlying assumptions of the radiative transfer model in FITSPEC. The program assumes a single continuum source behind all absorbing material and a constant temperature characterizing the material in each of up to four absorbing “clouds” along the line of sight. We also add an optically-thin emitting cloud, but neglect re-emission by the absorbing gas and absorption by the emitting gas. The column densities we report do not represent the entire column down to the embedded luminosity source, but rather the column only to an effective “dust photosphere” interior to which the continuum optical depth is large at a given wavelength. Consequently, our observations at $5 \mu\text{m}$ and $8 \mu\text{m}$ likely probe to greater depth than those at $13 \mu\text{m}$. The optical depth may be higher at longer wavelengths due to silicate opacity, and the colder, outer layers of the dusty region are likely too cool to emit significantly at $5 \mu\text{m}$. Comparison of results at different wavelengths is therefore difficult. We also

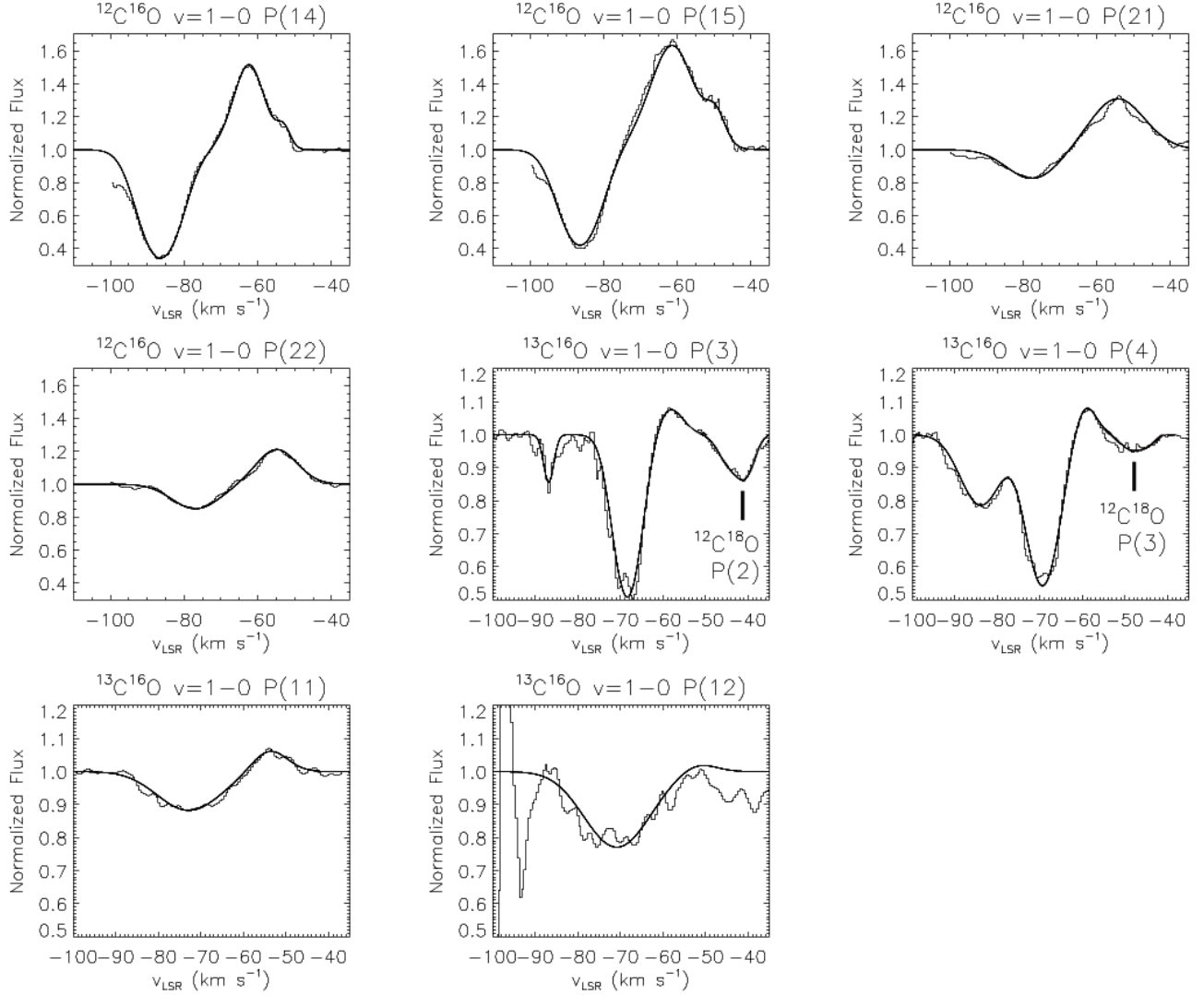


FIG. 12.— TEXES spectra of NGC 7538 IRS 9 (continued). Features shown are $^{12}\text{C}^{16}\text{O } v=1-0$ $P(14)$, $P(15)$, $P(21)$ and $P(22)$; $^{13}\text{C}^{16}\text{O } v=1-0$ $P(3)$ / $^{12}\text{C}^{18}\text{O } v=1-0$ $P(2)$; $^{13}\text{C}^{16}\text{O } v=1-0$ $P(4)$ / $^{12}\text{C}^{18}\text{O } v=1-0$ $P(3)$; and $^{13}\text{C}^{16}\text{O } v=1-0$ $P(11)$ and $P(12)$.

did not take into account emission by molecules following absorption of photons, except as approximated by our “negative absorption” method of treating emission lines. Furthermore, we ignored emission following radiative excitation. We assumed either pure absorption or optically-thin emission. However, the neglect of emission from a radiatively excited molecule should not be valid; at any plausible gas density the collisional de-excitation rate should be far below the radiative rate. Critical densities for vibrational transitions are typically greater than 10^{12} cm^{-3} , and almost all absorption events should be followed by emission, not collisional de-excitation. For two of our molecules, NH_3 and C_2H_2 , we were able to observe lines of more than one vibrational band or rotational branch, and the comparison of the observed line depths provides evidence that re-emission following absorption is an important effect.

For NH_3 , we observed weak emission in P -branch lines at wavelengths near $12 \mu\text{m}$ and possibly weak absorption in Q -branch lines near $10.5 \mu\text{m}$. R -branch lines were not observed due to the silicate and telluric O_3 absorption at their wavelengths. We explain the differences

between the P - and Q -branch lines based on the differing continuum radiation fields due to their differing proximity to the silicate dust feature. The NH_3 R branch is centered on the $9.7 \mu\text{m}$ silicate feature, the Q branch falls on its long wavelength shoulder, and the P branch is mostly between the 9.7 and $18 \mu\text{m}$ features. If the radiation field seen by a molecule is due to silicate emission either because the molecule is exposed to radiation from an optically thin external dust cloud or because it is immersed in a cloud of warm silicate dust, there will be more upward radiative transitions (i.e. more absorption) in the Q and R branches than in the P branch. The branching ratio between P -, Q -, and R -branch transitions from a given vibrationally-excited state depends on the K rotational quantum number, but is similar in the three bands. Consequently, if all absorption events are followed by an emission, absorptions will outnumber emissions in the R branch and emissions will outnumber absorptions in the P branch, resulting in net emission in the P -branch lines, as we observe. In Section 5.3 we describe a simple model involving purely radiative transitions in NH_3 incorporating upward and downward line

TABLE 3
MOLECULAR ABSORPTION PARAMETERS DETERMINED FROM χ^2
MINIMIZATION SPECTRAL FITTING

Species	v_{LSR} (km s ⁻¹)	T_{ex} (K)	b (km s ⁻¹)	$N(X)^a$ (10 ¹⁶ cm ⁻²)	C %
IRS 9 (this work)					
C ₂ H ₂	-58.9(1.4)	195(57)	4.0(1.3)	-0.2(0.1)	52(189) ^b
	-63.9(1.0)	205(49)	4.0(1.1)	2.3(0.7)	16(9)
HCN	-58.7(1.7)	200(68)	4.0(1.3)	0.5(0.7)	52(193) ^b
	-63.8(2.9)	204(76)	4.0(1.4)	3.9(4.3)	11(6)
NH ₃ ^a	-57.3(0.5)	365(47)	5.1(0.7)	-1.1(0.1)	52(195) ^b
	-61.6(0.6)	201(43)	5.9(1.1)	4.8(0.8)	8(8)
HNCO ^c	-58.9(0.0)	195(0)	4.0(0.1)	0.1(-0.5)	52(0)
	-63.9(10.0)	205(50)	4.0(1.0)	0.9(-1.6)	16(10)
IRS 1 (Knez et al. 2009)					
C ₂ H ₂	-55.7(0.3)	225(20)	0.6(0.1)	3.0(0.6)	6(0.6)
	-59.4(0.3)	191(10)	0.6(0.1)	2.8(0.2)	24(1)
HCN	-56.3(0.3)	256(30)	0.8(0.1)	5.6(2.8)	8(1)
	-60.0(0.3)	456(127)	1.0(0.4)	1.3(1.2)	68(50)
NH ₃	-57.3(0.3)	278(31)	0.9(1.0)	5.2(1.5)	9(1)
	-60.1(0.3)	248(32)	0.4(0.1)	2.8(0.8)	16(4)
HNCO	-57.2(0.3)	319(27)	1.8(0.2)	0.4(0.3)	4(0.3)
	-60.2(0.2)	171(79)	0.7(0.4)	0.1(0.3)	97(11)

NOTE. — Quantities in parentheses are the 1σ uncertainties on the adjacent figures in the same units unless otherwise indicated.

^a Negative column densities indicate components in emission. Column densities for the emission components are $C_{em}\tau_{em}(\nu)$ converted to N and do not represent a real column densities.

^b Unable to be constrained in the fit, assuming lines are optically thin.

^c All HNCO parameters other than the column density were fixed equal to those of C₂H₂ because an insufficient number of HNCO lines were observed to constrain them.

TABLE 4
3 σ UPPER LIMITS TO COLUMN DENSITY OF MOLECULES IN
THE UPPER STATE FOR SELECTED NON-DETECTIONS

Species/Mode	Line	n^a	$N_{limit}(X)$ (cm ⁻²)
¹³ CH ₄ ν_4^b	$R(1)$	1	4.1×10^{16}
	$R(4)$	4	3.1×10^{16}
OCS ν_2	P branch	34	2.9×10^{15}
HCN ν_2	$R(15)$	1	1.8×10^{16}
	$R(22)$	1	8.5×10^{16}
	$R(23)$	1	1.1×10^{17}
C ₂ H ₂ ν_5	$R(12)$	1	8.5×10^{15}
	$R(21)$	1	9.4×10^{15}
	$R(22)$	1	3.5×10^{16}
C ₂ H ₂ $2\nu_5 - \nu_5$	Q branch	15	2.1×10^{17}

NOTE. — A uniform temperature of 200 K was assumed for all species.

^a Number of lines/components in the attempted fit.

^b The given values of $N(^{13}\text{CH}_4)$ are for the isotopologue only rather than the total CH₄ column density.

strengths and assuming that the incident radiation field is proportional to the strength of the silicate feature at the wavelengths of the various lines. The model is time-independent and assumes the NH₃ gas is optically thin. More detailed and quantitative models will follow in a future paper.

5.2. C₂H₂, HCN, HNCO, OCS, and CO

We find that excitation temperatures for various molecules are basically consistent toward IRS 9 and IRS 1 with the exception of the hotter of the two NH₃ components toward IRS 1. That value of T_{ex} has a rather large uncertainty so it may in fact lie closer to ~ 250 K. The column densities are also similar in the two sources, though again with particular exception in the case of HCN. Knez et al. (2009) had the benefit of more observed HCN lines, but we find a lower column toward IRS 9. Most striking is the disagreement between the two sources in terms of the intrinsic line widths; although the thermal broadening should be similar, turbulence or systematic velocity variations along the line of sight to IRS 9 are greater than toward IRS 1. At the least, we can say that the spectral lines observed in the direction of IRS 9 are generally less saturated than those seen toward IRS 1, and this has important consequences for the accuracy of the IRS 9 column densities we obtain.

An example in which saturation of lines becomes important is CO. Knez et al. (2009) derived fractional abundances relative to CO for a variety of molecules seen in absorption toward IRS 1, attempting to compensate for badly saturated ¹²CO lines by using the value of $N(^{13}\text{CO})$ in Mitchell et al. (1990) and adopting $^{12}\text{C}/^{13}\text{C} = 45$ on the basis of their observations of C₂H₂. The value of $N(^{13}\text{CO})$ for IRS 9 reported by Mitchell et al. was based on their measurement of ¹²CO lines and assumed the terrestrial ratio $^{12}\text{C}/^{13}\text{C} = 89$. Boogert et al. (2004) confirm the saturation of the ¹²CO lines observed toward IRS 9. From observations of lines of ¹³CO and ¹²C¹⁸O in upper J levels of 7-15 they derive an optical depth $\tau_{12} = 13 \pm 4$ assuming $^{12}\text{C}/^{13}\text{C} = 80$ from Boogert et al. (2002). However, they argue that reliable values for $N(^{12}\text{CO})$ can be obtained, since the ¹³CO lines remain optically thin; they quote a value $N(^{12}\text{CO}) = (3.2 \pm 1.0) \times 10^{18}$ cm⁻². While we may argue with their choice of $^{12}\text{C}/^{13}\text{C}$ and with their conclusion that the ¹³CO lines are optically thin, their value of the CO column density is probably more realistic for IRS 9 than that obtained from the lower-resolution observations of Mitchell et al. (1990). Abundance ratios for some of the molecules we observed with TEXES, relative to this column of CO, are shown in Table 5.

Similarly, an estimate of the abundance of various molecules with respect to H₂ can be made given a measurement of the quantity of dust along the line of sight and a value for the gas-to-dust ratio in a typical protostellar envelope. Knez et al. adopted an H₂ column density toward IRS 1 of 7.5×10^{22} cm⁻² based on the 9.7 μm optical depth measured by Willner et al. (1982) to calculate fractional abundances of other molecules. Using the same method, we derive $N(\text{H}_2) = 5.2 \times 10^{22}$ cm⁻² toward IRS 9 and present the resulting abundances in Table 5. The pattern of relative abundances differs between the objects according to both the molecular species and whether the abundance of a given species is referred to CO or to H₂. The latter observation indicates that either the abundance ratio of CO to H₂ or the gas-to-dust ratio varies between objects.

The assumptions made in calculating the values in Table 5 – that the CO column density is comparable along both lines of sight and that the dust distribution and gas-to-dust ratio are similar in both cases –

TABLE 5
ABUNDANCES IN IRS 9 AND IRS 1 WITH RESPECT TO CO AND H₂

Molecule	IRS 9 (this work)		IRS 1 (Knez et al. 2009)	
	$N(X)/N(\text{CO})$	$N(X)/N(\text{H}_2)$	$N(X)/N(\text{CO})$	$N(X)/N(\text{H}_2)$
C ₂ H ₂	7.7(-3)	3.3(-7)	5.7(-3)	7.7(-7)
HCN	1.4(-2)	5.8(-7)	6.9(-3)	9.2(-7)
NH ₃	1.9(-2)	7.9(-7)	3.6(-2)	4.8(-6)
CH ₄	9.9(-2)	4.2(-6)	3.6(-2)	4.8(-6)
HNCO	<3.0(-3)	<1.3(-7)	5.0(-4)	6.0(-6)

NOTE. — All values are given in the notation $A(B) = A \times 10^B$. The column densities toward IRS 1 adopted for CO and H₂ were $N(^{12}\text{CO}) = 1 \times 10^{19} \text{ cm}^{-2}$ (Knez et al. 2009) and $N(\text{H}_2) = 7.5 \times 10^{22} \text{ cm}^{-2}$ (Willner et al. 1982). For IRS 9 the ¹²CO column density used was $N(^{12}\text{CO}) = 3.2 \times 10^{18} \text{ cm}^{-2}$ (Boogert et al. 2002) and the H₂ column density was the same as that used by Knez et al. (2009).

may not be unreasonable given that both objects are along nearby sightlines toward the same molecular cloud. Given this caveat, it is worth noting that the fractional abundances of all molecules relative to both CO and H₂ are broadly consistent with each other, except for HNCO with the values given for IRS 9 properly unconstrained by FITSPEC because no HNCO lines were detected. C₂H₂, HCN, and NH₃ are generally less abundant in IRS 9 than IRS 1, although this is only a suggestion in light of the assumptions.

Some kinematic information about IRS 9 is available in the line shapes. In combination with the varying effective depths to which sightlines probe at different wavelengths, line shapes give a sense of the distribution of various molecules along the line of sight. The P Cygni line profiles of our CO observations, for example, clearly indicate an outflow known previously. CH₄ absorption toward IRS 9 was known previously as well, but the TEXES data reveal the structure of the lines in more detail than the Keck+NIRSPEC spectra of Boogert et al. (2004) and the IRTF+Irshell spectra of Lacy et al. (1991).

In contrast to the circumstellar disk model of IRS 1 presented in Knez et al. (2009) based on TEXES observations, we cannot conclusively establish the presence of a disk in IRS 9. However, our fitting program preferred relatively small covering factors, similar to those in the best-fit synthetic spectra for the lower velocity component of IRS 1. Knez et al. interpret this as the signature of a near edge-on disk, acting as a source of continuum radiation that is only partially covered by the absorbing material. Given the high-velocity outflow in IRS 9 reported by Mitchell & Hasegawa (1991), we might rather be observing an accreting system nearly face-on in which the outflow both sweeps out a region near the continuum source and entrains material into the flow that only partially covers the source. This picture would account for the methane observations of Boogert et al. (2004) in which cold, solid-phase methane is seen in the outer envelope, becomes depleted in the outflow, and is seen in the gas phase in the inner envelope, all along a single sightline. We develop this model further in Section 5.6.

Non-detections of molecules also furnish some interpretable information. Carbonyl sulfide (OCS) is a molecule that commonly occurs in ices along sightlines through molecular clouds and indicates the presence of sulfur-bearing precursor species such as H₂S and SO₂

(Ferrante et al. 2008). Gibb et al. (2004) detected solid-phase OCS toward IRS 9 with *ISO*, finding a column density of $(5.5 \pm 4.4) \times 10^{15} \text{ cm}^{-2}$, and Knez et al. (2009) speculate that gas-phase OCS should be observable toward IRS 9 as indicated in some chemical models. We find an upper limit of $N(\text{OCS}) \leq 2.9 \times 10^{15} \text{ cm}^{-2}$ in the gas phase, which is compatible with the Gibb et al. result to within errors; we therefore conclude that any OCS near IRS 9 is firmly sequestered as an ice. Tielens (2005) cites an OCS ice abundance of 0.1% relative to H₂O ice toward IRS 9 given a water ice column density of $1 \times 10^{19} \text{ cm}^{-2}$. That figure may be compared to a CO ice abundance of 10%. In interstellar ices, OCS is commonly found in association with hydrogenated species such as H₂O and CH₃OH (Palumbo et al. 1997; Dartois et al. 1999). Laboratory results reported by Collings et al. (2004) indicate that OCS in an H₂O mixture desorbs at ~ 150 K, implying that significant quantities of this molecule may be frozen out in cold envelopes. However, C₂H₂ is seen toward both IRS 1 and IRS 9, for example, and according to the same study desorbs at only a slightly lower temperature. The relative columns of OCS and C₂H₂ may then reflect real differences in their abundances rather than their desorption temperatures. Alternately, OCS might only be abundant in the colder, outer envelope of IRS 9, explaining its apparent absence in the gas phase. Knez et al. (2009) did not observe OCS in the gas phase toward IRS 1 and Gibb et al. (2004) only cite an upper limit of $1 \times 10^{15} \text{ cm}^{-2}$ for the column density in the solid phase. Other sulfur-bearing molecules desorb from grains at a relatively low temperature of ~ 70 K, well below the temperatures in both objects. It is difficult, then, to understand why high-*J* lines of CS were observed toward IRS 1 by Knez et al. (2009) when recent models by Wakelam et al. (2011) *underpredict* the abundance of CS relative to other S-bearing molecules at temperatures above 100 K. The IRS 1 CS gas-phase column density is comparable to the solid-phase column density of OCS toward IRS 9, so if IRS 1 began with a comparable OCS abundance it may indicate a reaction route preference for forming CS as opposed to, e.g., SO₂. However, Wakelam et al. (2004) do not associate OCS and CS but rather assume CS exists in the gas of the pre-evaporative phase. We did not search for CS or SO₂ toward IRS 9 on the basis of their presumed low abundances, and thus are unable to comment on the likelihood

of any OCS/CS association.

The occurrence of isocyanic acid (HNCO) may be tied to the presence of the cyanate ion (OCN^-), where OCN^- is formed on grains by irradiation of solid HNCO in ices by UV in the presence of NH_3 (van Broekhuizen et al. 2004). Once evaporated from grains, OCN^- may then undergo hydrogenation to become HNCO again. We speculate, however, that the hydrogenation of OCN^- takes place *before* desorption, since the charge an ion induces while on a grain would result in a relatively high desorption temperature. Knez et al. (2009) observed dozens of HNCO lines toward IRS 1 and infer a column density of $5.4 \times 10^{15} \text{ cm}^{-2}$; this and the value of other gas parameters from their spectra led them to conclude that gas-phase HNCO may originate from evaporation of grain mantles by way of OCN^- . A low column density compared to the value expected from the observed abundance of OCN^- is interpreted as evidence that only a fraction of the available OCN^- desorbed from grains becomes HNCO, the remainder forming other molecules in gas-phase chemistry. Interpretation of observations is difficult, however, as Knez et al. point out that most chemical models do not currently include HNCO in their networks. All of the HNCO lines in the various TEXES settings we observed were non-detections. Our fitting program preferred column densities of $(0.1 \pm 0.5) \times 10^{16} \text{ cm}^{-2}$ and $(0.9 \pm 1.6) \times 10^{16} \text{ cm}^{-2}$ in the two presumed absorbing units, although a clear lack of detected lines make these values limits at best. We cannot, therefore, confidently distinguish between scenarios in which HNCO has a very low abundance in all phases and those in which it is effectively locked up in ices and unavailable to spectroscopy at our wavelengths.

5.3. Radiative Excitation and Radiative Transfer

In order to explain the observed pattern of NH_3 absorption and emission in our spectra of IRS 9, we invoke a radiative transfer effect that emerged from our modeling efforts. We begin with gas in a shell of thickness t around the continuum source with an NH_3 density n_{NH_3} . The rotational levels (J, K) of the ground vibrational state are assumed to be populated thermally at temperature T . Electric dipole transitions to and from an upper vibrational level proceed according to the selection rule that $\Delta K = 0$, $\Delta J = \pm 1$ or 0. In ammonia there is also a doubling of each vibrational level into symmetric and antisymmetric (s, a) inversion states, and the requirement that $s \leftrightarrow a$. Transitions are labeled by the inversion symmetry of the lower level. This gives the permitted transitions $a/s R(J_u-1, K)$, $a/s Q(J_u, K)$ and $a/s P(J_u+1, K)$. We obtain absorption coefficients α for each of these transitions from GEISA.

The absorption optical depth corresponding to α is

$$\tau(w) = \alpha N(\text{NH}_3) \phi_w \quad (13)$$

in which $N(\text{NH}_3)$ is the radial column density through the shell and ϕ_w is the line shape function, taken to be a Gaussian. Molecules are excited to the upper state along all branches at a rate

$$r_{up} = \sum_{P,Q,R} \int \frac{L_w}{hcw} \tau(w) dw \quad (14)$$

per second assuming $\tau \ll 1$ through the shell, where L_w

is the incident continuum luminosity. The rate at which photons are emitted on a given branch is determined by the rate of transitions populating the upper level and the branching ratio of the Einstein A coefficients, e.g.,

$$r_{down}(P) = r_{up}(P) \frac{A_P}{A_P + A_Q + A_R} \quad (15)$$

for the P branch. We calculate values for the A coefficients according to

$$A_{ul} = \frac{8\pi c}{\lambda^2} \alpha(T) \frac{Q_r}{g_u} e^{E_l/kT} \quad (16)$$

where λ is the wavelength of the transition, Q_r is the rotational partition function, g_u is the statistical weight of the upper state, and E_l is the energy of the lower state.

We consider the $aP(4, K)$, $aQ(3, K)$, and $aR(2, K)$ lines that go to the $J = 3$ rotational level of the sv_2 vibrational level. We use the branching ratios to compute the expected emission flux in a given line l with wavenumber center w from

$$F_l = \frac{A_l}{A_P + A_Q + A_R} \frac{hcw}{4\pi d^2} \sum_{P,Q,R} \frac{L_w}{hcw} N \alpha(T). \quad (17)$$

To calculate the net flux in a line we also need the amount of absorption. We require the relative intensity of the $10 \mu\text{m}$ dust emission feature at the frequency of each of the branches in order to compute the absorption line flux. We use mid-infrared spectra of the young stars DI Cep and DK Tau obtained by Hanner et al. (1998) to which they fit models of optically-thin silicate dust emission. From these spectra we calculate relative intensities of 0.52, 0.78 and 1.00 for the continuum at the 11.7, 10.7, and $9.7 \mu\text{m}$ wavelengths of the aP , aQ , and aR lines for which $J_u = 3$, respectively. The intensities are normalized to the value on the aR branch as it is nearest to the peak of the $10 \mu\text{m}$ silicate feature. The observed continuum flux near the P -branch lines, $\sim 60 \text{ Jy}$, is turned into a luminosity at the source by multiplying by $4\pi d^2$, where d is the distance to NGC 7538 (2.65 kpc). We corrected for extinction using the $10 \mu\text{m}$ optical depth measurement of Willner et al. (1982), which implies $A_{11.7\mu\text{m}} \sim 2.2$ magnitudes and $A_{9.7\mu\text{m}} \sim 4.4$ magnitudes toward IRS 9.

The extinction-corrected luminosity $\sim 1.5 \times 10^{35} \text{ erg s}^{-1} \text{ cm}^{-1}$ at the position of the P -branch lines is scaled to the other branches according the values above. We assume throughout that the lines are sufficiently broad that they remain optically thin, leading to uncertainty in the covering factor. Only the product of column density and covering factor is meaningful under such circumstances, so we used the sum of this product for the two NH_3 components rather than the sum of the NH_3 column densities themselves in calculating F_l . We considered the possibility that the optically thin case might result in uncertainties propagating to the line fluxes, and to this end we carried out additional runs of FITSPEC in which we deliberately held the value of the covering factor fixed equal to 1. This resulted in a variation of the product of column density and covering factor of $< 20\%$ compared to allowing the covering factor to vary. We therefore retained the assumption of the optically thin case, and opted to use the lower value for the product of column density

TABLE 6
PREDICTED FLUXES FOR THE aP, aQ, aR
LINES OF ν_2 TRANSITIONS OF NH_3 (J_u
= 3) AND MEASURED FLUXES FOR THE
CORRESPONDING aP -BRANCH LINES
FROM TEXES SPECTRA.

K	Predicted			Observed F_{aP}
	F_{aP}	F_{aQ}	F_{aR}	
0	5.55	—	-1.94	1.80
1	2.64	0.03	-0.95	1.35
2	2.03	0.02	-0.73	—
3	2.11	-0.87	—	0.90

NOTE. — All fluxes are given in units of $10^{14} \text{ erg s}^{-1} \text{ cm}^{-2}$. Positive flux values indicate net emission while negative values indicate net absorption in a given line.

and covering factor of $8.52 \times 10^{15} \text{ cm}^{-2}$ in place of N in Equation 17. For $\alpha(T)$, we chose an excitation temperature of 200 K, consistent with the FITSPEC results for the -61.6 km s^{-1} component. Finally, the assumed extinction was re-introduced at the end of the calculation to realistically predict the line fluxes observed at Earth.

The line fluxes we compute are given in Table 6. The predicted fluxes are the differences of the strengths of the predicted emission and absorption components of each line and show an overall pattern of emission on the P branch, weak absorption or no lines on the Q branch, and absorption on the R branch. The P -branch fluxes compare favorably with the integrated line fluxes from our data implying that our explanation is at least plausible. This suggests that the dust and NH_3 are in close proximity to one another, if not actually mixed.

In the case of acetylene, we observed lines of the ν_5 Q and R branches at $12.5\text{--}13.7 \mu\text{m}$ and the $\nu_4 + \nu_5$ P branch near $7.5 \mu\text{m}$. The ν_5 Q -branch and $\nu_4 + \nu_5$ P -branch lines are of low S/N, due to telluric interference, but clearly lead to higher C_2H_2 column densities than the ν_5 R -branch lines when the observations are interpreted with the pure absorption model. The $\nu_4 + \nu_5$ lines appear to require at least an order of magnitude more C_2H_2 than do the ν_5 R -branch lines. The radiative transfer effects discussed for NH_3 may also be relevant for C_2H_2 , and may explain a part of the discrepancy between the C_2H_2 bands. But the bands are farther from the silicate feature and the wavelength separation between the branches of the C_2H_2 bands is smaller than for NH_3 , so the effects should be smaller. The greater column density derived from the $\nu_4 + \nu_5$ lines could instead be due to the smaller extinction at $7.5 \mu\text{m}$ allowing a greater column of gas to be probed. This explanation was suggested by Evans et al. (1991) to explain a similar effect seen toward Orion IRC2, but it is difficult to understand how it could explain more than a factor ~ 2 discrepancy, given the ratio of extinction at the different wavelengths.

We propose a different explanation for the variation between the column densities derived from the C_2H_2 ν_5 and $\nu_4 + \nu_5$ lines. Vibrational excitation caused by a ν_5 band photon is almost always followed by emission of another ν_5 band photon, although possibly in a different branch. (Due to the C_2H_2 selection rules, P - or R -branch absorption can be followed by either P - or R -branch emission,

but Q -branch absorption must be followed by emission in the same branch.) However, $7.5 \mu\text{m}$ $\nu_4 + \nu_5$ absorption is most often followed by a $\nu_4 + \nu_5 - \nu_4$ transition at $\sim 13.7 \mu\text{m}$ to the ν_4 level. Since radiative decay from the ν_4 level is forbidden, the ν_4 excitation energy is then eventually lost in a collision. The result of this series of transitions is that there is little $\nu_4 + \nu_5$ emission canceling the $\nu_4 + \nu_5$ absorption, and these lines are probably reasonably well treated with the pure absorption model that we used. In contrast, the ν_5 absorption lines may be largely cancelled by emission, leading to an underestimate of the C_2H_2 column density. The amount of re-emission into our beam is very dependent on the (unknown) source geometry, which determines whether our line of sight has more or less than the average for the gas surrounding the source. Assuming the $\nu_4 + \nu_5$ lines can be treated with a pure absorption model to give a fair measure of the abundance along our line of sight, the required correction to the column densities derived from the ν_5 lines is approximately a factor of 10. The required correction to the column densities of our other observed molecules may be similar.

5.4. Abundance Trends and Relative Ages

In Figure 13 we present a graphical comparison of the column densities of C_2H_2 , HCN , NH_3 , HNCO , and CH_4 we obtain toward IRS 9 with those derived for IRS 1 in Knez et al. (2009). Column densities for each of the two components with similar LSR velocities (-57 km s^{-1} and -61 km s^{-1}) toward each object are shown. We reach rather different conclusions for the two velocity components; IRS 1 appears to be clearly enhanced in all molecules except for HNCO at -57 km s^{-1} , whereas the column densities are essentially the same to within errors at -61 km s^{-1} . HNCO is the exception at the higher negative velocity, but we again stress that the cited values for IRS 9 are those preferred by our fitting program despite HNCO being an evident non-detection in our data. It is therefore not strictly reliable in comparison with the firm detection of HNCO toward IRS 1. Also, we do not have any evidence *a priori* that these two velocity components are kinematically related. Knez et al. (2009) claim the two components seen toward IRS 1 are real but may not be separate; they may instead represent a non-Gaussian velocity distribution. If the latter situation is true, then such a velocity distribution may be systemic in other parts of NGC 7538. Regardless, the pattern that emerges from Figure 13 is suggestive of real differences in the gas-phase abundance of these molecules, which we attribute to differing thermal conditions in the environment of each object. If the molecular abundances in all phases were the same to begin with, the gas-phase abundances we see at the present time are indicative of a difference in the relative ages of IRS 9 and IRS 1.

Several other lines of evidence also argue in favor of a younger age for IRS 9. First, a number of molecules (e.g., H_2O , CH_3OH , HCOOH , and OCN^-) are seen in the solid phase in the spectrum of IRS 9 that are either in the gas phase or apparently absent toward IRS 1. CO and CH_4 are seen both in the solid and gas phases, toward IRS 9 although the gas-phase lines probably form in the shocked outflow material. Second, a fairly strong ionizing radiation field exists near IRS 1, whereas our

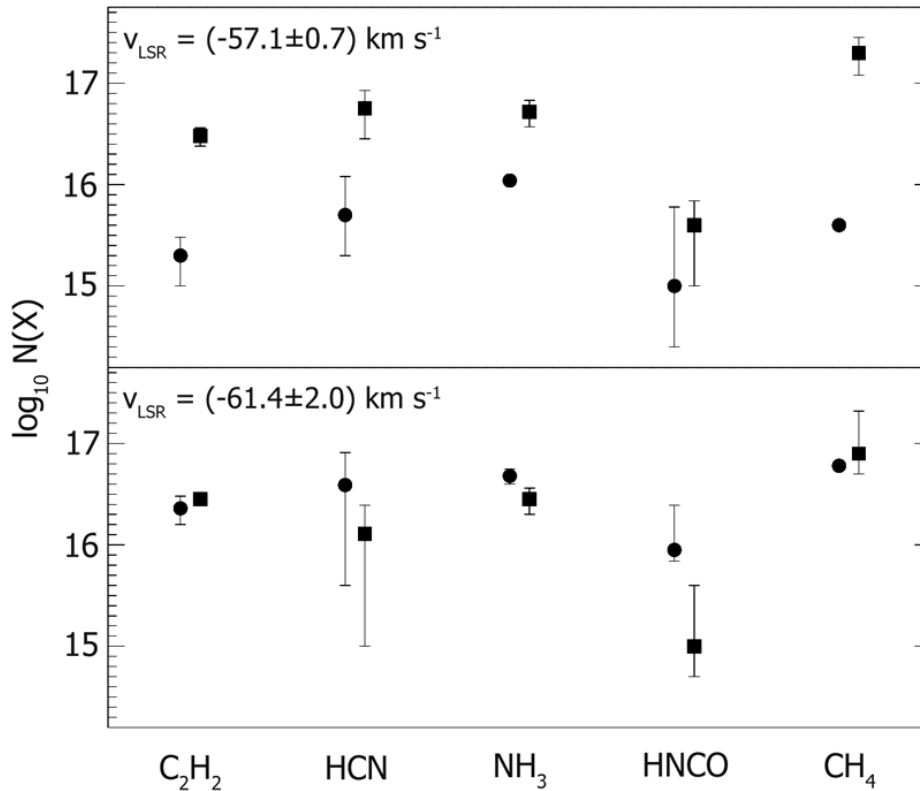


FIG. 13.— Log column densities for the main molecular species detected in this study for the two principal velocity components toward each target. IRS 9 values (this work) are shown as filled circles and IRS 1 values from Knez et al. (2009) as filled squares. The symbols for a given molecule are slightly offset for clarity.

non-detection of the $12.8 \mu\text{m}$ line of [Ne II] discussed in Section 4.4 argues for a much weaker field toward IRS 9. In addition, IRS 1 has an ionized component to its outflow, probably driven by a wind; IRS 9 shows no similar outflow component. Third, the chemistry of IRS 1 resembles the “hot core” phase in which ices desorb from dust grains, leading to a short-lived ($\sim 10^4$ – 10^5 yr) high-temperature gas-phase chemistry and the synthesis of complex organic molecular species (van der Tak 2005). Some species, such as HNCO, are clear detections toward IRS 1 but appear to be absent toward IRS 9; we cannot rule out an intrinsic abundance effect in explaining this difference, but a simpler explanation is that the IRS 9 core is not yet sufficiently warm to evaporate it.

Knez et al. (2009) claim an age of $\sim 10^5$ yr for IRS 1 based on their chemical models. However, there is significant uncertainty in that value. Abundances of the molecules they detected indicate an age closer to $\sim 10^4$ yr, but for other molecules the chemical model predictions disagree with the observed values, suggesting ages in the range of 2×10^3 to 2×10^6 yr. They explain this large range as a result of insufficiently detailed models. The models of Doty et al. (2002) predict enhancements in the abundances of C₂H₂, HCN, and CH₄ for temperatures above about 800 K at late times ($\geq 10^5$ yr); the enhancements are similar to those observed toward IRS 1, but the temperatures derived by Knez et al. (2009) are considerably lower than this. The fractional abundance of HCN in the Doty et al. model for 200 K quickly turns down at times $> 10^5$ yr whereas those of C₂H₂ and CH₄ increase through at least $10^{5.5}$ yr; $N(\text{HCN})$ is in fact

slightly higher toward IRS 1 compared to IRS 9 by a factor of about 1.4. This may both support the argument that IRS 1 is the older object of the two and suggest an upper limit for its age, since the HCN abundance has not begun the predicted decline, but this conclusion is tentative since our value of $N(\text{HCN})$ for IRS 9 is determined by the singular detection of the ν_2 $R(16)$ line. Lastly, an additional age constraint from chemistry is provided by the non-detection of the molecule CCS by Hoffman & Seojin Kim (2011) in EVLA data, suggesting an age $\geq 10^4$ yr. Given that we find IRS 9 to be in an evidently more primitive chemical state than IRS 1, the upper limit to its age is probably $\sim 10^4$ yr. Detailed, robust chemical modeling, using our observations as inputs, may help more firmly constrain this value.

Our results, when considered in the context of the earlier work on IRS 1, qualitatively support a picture of an evolutionary sequence of cores in NGC 7538 consistent with the models of Elmegreen & Lada (1977) and Campbell & Thompson (1984) in which a wave of shock/ ionization fronts from a previous generation of stars caused the sequential collapse of the cores that became IRS 1 and IRS 9, respectively. As the wave continues to propagate toward the southeast, it should initiate the collapse of cores found in progressively more primitive states. The ‘IR-quiet’ protostellar object NGC 7538 S is likely the next youngest object in the region after IRS 9 (Pestalozzi et al. 2006). Such a sequence, including NGC 7538 IRS 11 among the youngest objects, is supported by the maser observations of Hutawarakorn & Cohen (2003). It is tempting to assume that the wave has left the IRS 1

region already, but several newly-discovered submillimeter sources within 0.35 pc of IRS 1 reported by Qiu et al. (2011), presumably harboring forming intermediate- or high-mass stars, may be younger than IRS 9 and contemporaneous with NGC 7538 S. However, they may also be fragments, along with IRS 1, of a common progenitor. The star formation history of NGC 7538 may therefore be more complex than can be represented with a simple, monotonically-increasing series of protostellar ages.

5.5. Spatial Resolution of IRS 9

We checked whether IRS 9 is spatially resolved in TEXES data by obtaining spatial-spectral maps of IRS 9 and μ Cep, a mid-infrared bright standard star that was used as a PSF reference for spatial deconvolution. To generate the maps, the TEXES slit was stepped across each object and spectral information gathered at each step, building up data cubes in which there are two spatial dimensions and one spectral dimension. Additional steps were included at both the beginning and end of the scans to sample the sky for removal. The data are handled by our reduction pipeline in a manner substantially similar to the method of processing the nodded observations, treating a scan as a series of 2-D spectrograms.

IRS 9 and μ Cep were observed with identical scan parameters including scan direction, length, and step size, and were observed at comparable zenith angles. After reducing the raw data, different spectral windows were selected from which to form reconstructed images of each object by summing the data in the spectral direction. For IRS 9, we chose a window centered at about 745.88 cm^{-1} , near the position of the $\text{C}_2\text{H}_2 \nu_5 R(6)$ line in the rest frame of the object. The window was 0.35 cm^{-1} wide. We selected this frequency range to search for spatial extent in the IRS 9 map as we noted absorption in the $R(6)$ line in nod-mode spectra. For the deconvolution reference, we chose a nearby spectral region apparently free from telluric and stellar photospheric absorption centered at 744.43 cm^{-1} and 0.25 cm^{-1} in width. The observed FWHM for each object was $0''.67$ for IRS 9 and $0''.63$ for μ Cep at each object's respective spectral setting. The deconvolution was performed using the maximum entropy method (Bryan & Skilling 1980; Narayan & Nityananda 1986), with results shown in Figure 14. The deconvolved image of IRS 9 in the C_2H_2 feature appears pointlike with some possible weak extension toward the NE. We do not believe this indicates that the envelope or other structure of IRS 9 is spatially resolved, and we carried out our other TEXES observations exclusively in nod mode.

This approach to imaging and deconvolution should be interpreted cautiously. Poor seeing during a scan, for example, might distort the PSF on time scales shorter than the time required to scan across an object. However, we felt sufficiently convinced that IRS 9 presents itself as a point source in the mid-infrared that we did no further scans. Given the distance to NGC 7538 and the spatial resolution of our observations, we conclude that spatial structure of IRS 9 is unresolved on scales $\leq 2000 \text{ AU}$. This limit is consistent with the inferred sizes of disks detected around protostars with $10^3 - 10^5 L_\odot$ (Patel et al. 2005; Schreyer et al. 2006; Rodríguez et al. 2007; Franco-Hernández et al. 2009; Galván-Madrid et al. 2010) and places a constraint on physical models of mas-

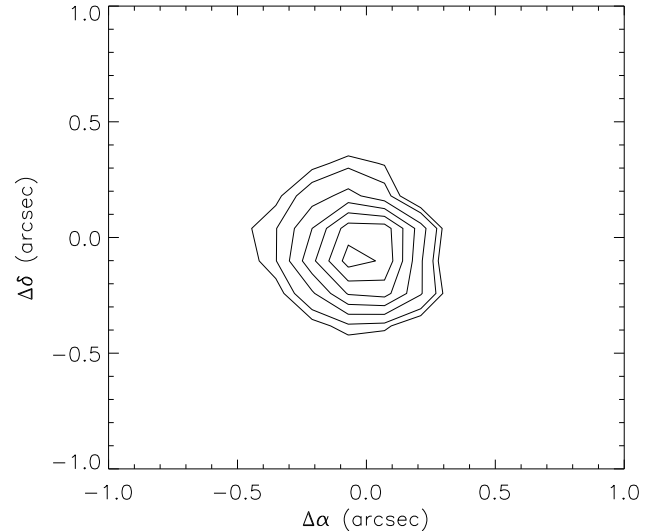


FIG. 14.— Contour plot of the TEXES scan map of NGC 7538 IRS 9 at 745.8 cm^{-1} after maximum entropy deconvolution. The PSF reference for the deconvolution was a scan map of the mid-infrared standard μ Cep made with the same parameters as the IRS 9 scan. North is up and east at left and contours are plotted at flux densities of 25, 50, 100, 150, 200, 250 and 300 Jy . The plot origin is approximately the location of the source peak intensity.

sive protostars in NGC 7538, which often include disks thought to be hundreds of AU in size (e.g., Sandell et al. 2009; Surcis et al. 2011).

5.6. A Structural Model Of IRS 9

Based on the conclusions we reach from kinematic indications in the data, we compose a toy model of the IRS 9 system to show its essential structure and a guess at the viewing geometry. This model is shown in Figure 15. In this view, the central, perhaps accreting object is shown in cross-section with a surrounding disk. The known bipolar outflow is indicated by the evacuated regions above and below the disk and by arrows showing the direction of the flow. A lighter shading above and below the disk represents the remaining envelope that has not been accreted or become part of the outflow. In order to account for seeing certain molecules under presumed LTE conditions while not observing the exciting source directly, we suggest a “dust photosphere” at a radius from the source at which the emission from the dust becomes optically thick, as indicated by dashed lines in the figure. It is this effective source of radiation against which we see most of our absorption lines. Since we do not resolve spatial structure in the molecular absorption region above the dust photosphere, we interpret the size scale constraint discussed in the previous section (2000 AU) as an upper limit to the radius of the dust photosphere itself. This value also represents an upper limit of the size of the disk, because a larger, optically-thick disk should have been detectable by virtue of its spatial extent. Our data do not permit a constraint to be placed on the size of the surrounding envelope or the outflow cavities.

The inclusion of a disk in this picture is tentative, given that the lifetime of disks around massive stars is thought to be particularly short, of order 10^{-4} the lifetime of the

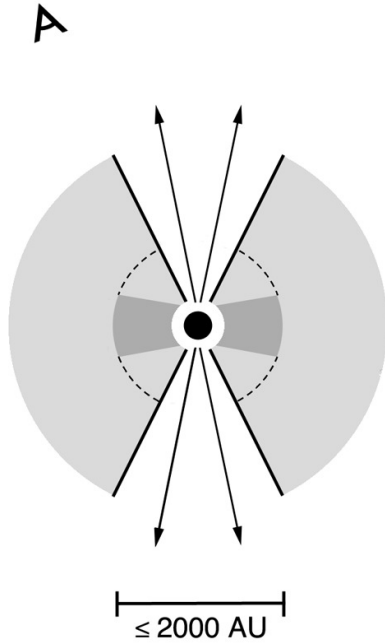


FIG. 15.— A cartoon rendering for a proposed model of NGC 7538 IRS 9 based on TEXES data. A central, embedded source (or sources) has at least partially cleared cavities in a bipolar outflow while significant dust remains in a spherical distribution (light gray shading). An inner disk may exist (dark gray shading). To account for our observations, the line of sight samples both scattered light in the outflow as well as a “dust photosphere” (dashed lines) where the dust becomes opaque at a mid-infrared wavelengths. The inferred viewing angle is shown, and a scale bar is provided based on the constraints discussed in Section 5.5. Note that the envelope and outflow cavities as depicted are not shown to this scale.

host star (Mann & Williams 2009). The evolutionary state of IRS 9 may be sufficiently early in its history, well before the embedded protostar is visible, to retain a disk. Knez et al. (2009) argue for a disk around IRS 1 whose surface is observed at near-grazing incidence; a sightline through the disk atmosphere is proposed to explain their observations.

We propose a viewing angle based on the following. First, the broad CO lines with their large, Doppler-shifted velocity separations suggest formation in the bipolar outflow. Second, the apparent radiative “pumping” of the NH_3 transitions, discussed in Section 5.3, requires continuum light from the direction of the dust photosphere. Third, the other molecular species we observed in absorption also require a background continuum source, though the range of excitation conditions we derive is insufficient to definitively locate them in the envelope or the turbulent edge of the outflow cavity. The C_2H_2 combination mode may be an exception and could shock-excited in the outflow. Consequently, we place the viewing angle such that the observer is looking nearly along the edge of the outflow cavity and sees light from both sources. This model appears to account for all aspects of our observations in at least a qualitative way. It is also consistent with the conclusion of Sandell et al. (2005), based on $\text{HCO}^+ J = 2 \rightarrow 1$ observations showing deep, redshifted self-absorption, that the molecular outflow is viewed nearly pole-on.

6. SUMMARY

We have presented the method and results of a study of the embedded high-mass protostellar object NGC 7538 IRS 9 and compared our findings to a similar investigation of the related object NGC 7538 IRS 1. We obtained high resolution, mid-infrared spectra of these objects in 46 ro-vibrational transitions of the fundamental bands of the molecules C_2H_2 , CH_4 , HCN , NH_3 and CO and a number of their isotopologues. We also detected two lines of the $\nu_4 + \nu_5$ combination mode of C_2H_2 . From these observations we draw some broad conclusions.

1. IRS 9 appears to be a spatially-unresolved object on a scale of ~ 2000 AU.

2. With the exception of CO , whose lines are saturated, we did not observe ^{13}C isotopologues of any organic molecules. This may indicate a particularly large value of the $^{12}\text{C}/^{13}\text{C}$ ratio, or the ^{13}C isotopologue lines were simply below our threshold of detection.

3. There is no discernible, consistent trend in the abundance variations of C_2H_2 , HCN , NH_3 , and CH_4 with respect to either CO or H_2 between IRS 9 and IRS 1. Column density variations between the two objects are an order of magnitude or less in each case.

4. The observation of gas-phase CS absorption toward IRS 1 and the evident non-detection of sulfur-bearing species like OCS in the gas phase toward IRS 9 may reflect a real S abundance variation between the two objects but can also be explained by different thermal conditions in each case. A similar abundance discrepancy involving HNC may indicate sufficiently low temperatures in the IRS 9 envelope to retain it as an ice and explain its detection in the solid phase by Gibb et al. (2004). If the abundances in all phases are actually comparable between the two sources, a relatively weak heating source is implied in the case of IRS 9. While this is at odds with the somewhat higher inferred abundances of the other molecules toward the warmer environment of IRS 1, the suggested sightline in that direction through a disk atmosphere subjects those molecules to direct irradiation by protostellar UV. This may drive photochemistry that lowers abundances of certain molecules as they are converted to other kinds.

5. However, the relative abundances of various molecular species with respect to CO and H_2 can also be explained by variations in the CO/H_2 ratio or the gas-to-dust ratio between the two objects. Choosing between these interpretations would require better determining the CO and H_2 abundances toward both objects.

6. Lines of many species observed toward IRS 9 have higher Doppler b values than those toward IRS 1, resulting in lines toward IRS 9 being less saturated than those toward IRS 1.

7. Given the non-detection of Ne II in our data and the continuum flux density at the expected wavenumber of the line (70 ± 10 Jy), we find a corresponding upper limit of $\sim 10^{45}$ photons s^{-1} for the Lyman continuum flux in IRS 9. This value is well below both the expected flux of a ZAMS star with the same luminosity and the ionizing flux of IRS 1 ($> 10^{48}$ photons s^{-1}).

8. The observed pattern of emission and absorption components of ammonia lines in our spectra can be explained by a non-LTE radiative transfer effect involving “pumping” of lines on the R -branch by light from the $9.7 \mu\text{m}$ silicate dust feature. Predictions of the strength of

this effect compare reasonably well with observed fluxes in *P*-branch lines.

9. Our data are consistent with a simple model of the IRS 9 system in which the sightline probes both a high-speed outflow and a quiescent envelope illuminated from behind by a “dust photosphere” such that line formation depth is wavelength-dependent. The possibility of seeing the turbulent flow along the wall of the outflow cavity is supported by the broad, blueshifted wing of the C_2H_2 $\nu_4 + \nu_5$ combination band that may arise in a shock. A disk is included in the model given the system’s inferred young age and the presence of a kinematically-indicated disk toward the (presumably) older IRS 1.

10. Observed differences in the abundance of various molecular species toward IRS 9 and IRS 1 imply an upper limit to the age of IRS 9 of $\sim 10^4$ yr, whereas IRS 1 may be an order of magnitude older. This conclusion is consistent with the triggered star formation models of Elmegreen & Lada (1977) as applied to NGC 7538 by Campbell & Thompson (1984) and others.

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REFERENCES

- Akabane, K., Matsuo, H., Kuno, N., & Sugitani, K. 2001, *PASJ*, 53, 821
- Araya, E., Hofner, P., Goss, W. M., et al. 2007, *ApJS*, 170, 152
- Arnett, D. 1996, *Supernovae and Nucleosynthesis: An Investigation of the History of Matter from the Big Bang to the Present* (Supernovae and Nucleosynthesis: An Investigation of the History of Matter, from the Big Bang to the Present, by D. Arnett. Princeton: Princeton University Press, 1996.)
- Bally, J. & Zinnecker, H. 2005, *AJ*, 129, 2281
- Balog, Z., Kenyon, S. J., Lada, E. A., et al. 2004, *AJ*, 128, 2942
- Beichman, C. A., Neugebauer, G., Habing, H. J., Clegg, P. E., & Chester, T. J., eds. 1988, *Infrared astronomical satellite (IRAS) catalogs and atlases. Volume 1: Explanatory supplement, Vol. 1*
- Beuther, H., Linz, H., & Henning, T. 2012, *A&A*, 543, A88
- Bevington, P. R. & Robinson, D. K. 2003, *Data reduction and error analysis for the physical sciences* (Data reduction and error analysis for the physical sciences, 3rd ed., by Philip R. Bevington, and Keith D. Robinson. Boston, MA: McGraw-Hill, ISBN 0-07-247227-8, 2003.)
- Blaauw, A. 1964, *ARA&A*, 2, 213
- Blaauw, A. 1991, in *NATO ASIC Proc. 342: The Physics of Star Formation and Early Stellar Evolution*, ed. C. J. Lada & N. D. Kylafis, 125
- Bonnell, I. A. & Bate, M. R. 2005, *MNRAS*, 362, 915
- Bonnell, I. A., Bate, M. R., Clarke, C. J., & Pringle, J. E. 2001, *MNRAS*, 323, 785
- Bonnell, I. A., Bate, M. R., & Zinnecker, H. 1998, *MNRAS*, 298, 93
- Bonnell, I. A., Vine, S. G., & Bate, M. R. 2004, *MNRAS*, 349, 735
- Boogert, A. C. A., Blake, G. A., & Öberg, K. 2004, *ApJ*, 615, 344
- Boogert, A. C. A., Blake, G. A., & Tielens, A. G. G. M. 2002, *ApJ*, 577, 271
- Brunt, C. M. 2003, *ApJ*, 584, 293
- Bryan, R. K. & Skilling, J. 1980, *MNRAS*, 191, 69
- Campbell, B. & Thompson, R. I. 1984, *ApJ*, 279, 650
- Collings, M. P., Anderson, M. A., Chen, R., et al. 2004, *MNRAS*, 354, 1133
- Dartois, E., Demyk, K., d’Hendecourt, L., & Ehrenfreund, P. 1999, *A&A*, 351, 1066
- Dickel, H. R., Dickel, J. R., & Wilson, W. J. 1981, *ApJ*, 250, L43
- Doty, S. D., Schöier, F. L., & van Dishoeck, E. F. 2004, *A&A*, 418, 1021
- Doty, S. D., van Dishoeck, E. F., van der Tak, F. F. S., & Boonman, A. M. S. 2002, *A&A*, 389, 446
- Ekström, S., Georgy, C., Eggenberger, P., et al. 2012, *A&A*, 537, A146
- Elmegreen, B. G. & Lada, C. J. 1977, *ApJ*, 214, 725
- Evans, II, N. J., Lacy, J. H., & Carr, J. S. 1991, *ApJ*, 383, 674
- Ferrante, R. F., Moore, M. H., Spiliotis, M. M., & Hudson, R. L. 2008, *ApJ*, 684, 1210
- Fischer, J., Joyce, R. R., Righini-Cohen, G., Simon, M., & Simon, T. 1980, *ApJ*, 240, L95
- Foster, T. & Routledge, D. 2003, *ApJ*, 598, 1005
- Franco-Hernández, R., Moran, J. M., Rodríguez, L. F., & Garay, G. 2009, *ApJ*, 701, 974
- Galván-Madrid, R., Zhang, Q., Keto, E., et al. 2010, *ApJ*, 725, 17
- Gibb, E. L., Whittet, D. C. B., Boogert, A. C. A., & Tielens, A. G. G. M. 2004, *ApJS*, 151, 35
- Hackwell, J. A., Grasdalén, G. L., & Gehr, R. D. 1982, *ApJ*, 252, 250
- Hanner, M. S., Brooke, T. Y., & Tokunaga, A. T. 1998, *ApJ*, 502, 871
- Henkel, C., Wilson, T. L., & Johnston, K. J. 1984, *ApJ*, 282, L93
- Heyer, M. H., Carpenter, J. M., & Snell, R. L. 2001, *ApJ*, 551, 852
- Ho, L. C. & Keto, E. 2007, *ApJ*, 658, 314
- Hoffman, I. M. & Seojin Kim, S. 2011, *AJ*, 142, 202
- Hutawarakorn, B. & Cohen, R. J. 2003, *MNRAS*, 345, 175
- Jacquinet-Husson, N., Scott, N. A., Chédin, A., et al. 2005, *Journal of Quantitative Spectroscopy and Radiative Transfer*, 95, 429
- Jiang, Z., Tamura, M., Hoare, M. G., et al. 2008, *ApJ*, 673, L175
- Johnston, K. G., Keto, E., Robitaille, T. P., & Wood, K. 2011, *MNRAS*, 415, 2953
- Kameya, O. 2006, in *Proceedings of the 8th European VLBI Network Symposium*
- Keto, E. & Zhang, Q. 2010, *MNRAS*, 406, 102
- Knez, C., Lacy, J. H., Evans, N. J., van Dishoeck, E. F., & Richter, M. J. 2009, *ApJ*, 696, 471
- Lacy, J. H., Carr, J. S., Evans, II, N. J., et al. 1991, *ApJ*, 376, 556
- Lacy, J. H. & Jaffe, D. T. 2007, in *American Astronomical Society Meeting Abstracts, Vol. 211, American Astronomical Society Meeting Abstracts, #76.03*
- Lacy, J. H., Richter, M. J., Greathouse, T. K., Jaffe, D. T., & Zhu, Q. 2002, *PASP*, 114, 153
- Lugo, J., Lizano, S., & Garay, G. 2004, *ApJ*, 614, 807
- Mann, R. K. & Williams, J. P. 2009, *ApJ*, 694, L36
- McKee, C. F. & Ostriker, E. C. 2007, *ARA&A*, 45, 565
- Michelson, A. A. 1898, *ApJ*, 8, 37
- Mitchell, G. F. & Hasegawa, T. I. 1991, *ApJ*, 371, L33

- Mitchell, G. F., Maillard, J.-P., Allen, M., Beer, R., & Belcourt, K. 1990, *ApJ*, 363, 554
- Moscadelli, L., Reid, M. J., Menten, K. M., et al. 2009, *ApJ*, 693, 406
- Narayan, R. & Nityananda, R. 1986, *ARA&A*, 24, 127
- Osterbrock, D. E. & Ferland, G. J. 2006, *Astrophysics of gaseous nebulae and active galactic nuclei* (*Astrophysics of gaseous nebulae and active galactic nuclei*, 2nd. ed. by D.E. Osterbrock and G.J. Ferland. Sausalito, CA: University Science Books, 2006)
- Palumbo, M. E., Geballe, T. R., & Tielens, A. G. G. M. 1997, *ApJ*, 479, 839
- Panagia, N. 1973, *AJ*, 78, 929
- Patel, N. A., Curiel, S., Sridharan, T. K., et al. 2005, *Nature*, 437, 109
- Pendleton, Y. J., Tielens, A. G. G. M., Tokunaga, A. T., & Bernstein, M. P. 1999, *ApJ*, 513, 294
- Pestalozzi, M. R., Elitzur, M., Conway, J. E., & Booth, R. S. 2004, *ApJ*, 603, L113
- Pestalozzi, M. R., Minier, V., Motte, F., & Conway, J. E. 2006, *A&A*, 448, L57
- Qiu, K., Zhang, Q., & Menten, K. M. 2011, *ApJ*, 728, 6
- Reid, M. A. & Wilson, C. D. 2005, *ApJ*, 625, 891
- Rodríguez, L. F., Zapata, L. A., & Ho, P. T. P. 2007, *ApJ*, 654, L143
- Sandell, G., Goss, W. M., & Wright, M. 2005, *ApJ*, 621, 839
- Sandell, G., Goss, W. M., Wright, M., & Corder, S. 2009, *ApJ*, 699, L31
- Schreyer, K., Semenov, D., Henning, T., & Forbrich, J. 2006, *ApJ*, 637, L129
- Smith, R. L., Pontoppidan, K. M., Young, E. D., Morris, M. R., & van Dishoeck, E. F. 2009, *ApJ*, 701, 163
- Surcis, G., Vlemmings, W. H. T., Torres, R. M., van Langevelde, H. J., & Hutawarakorn Kramer, B. 2011, *A&A*, 533, A47
- Tielens, A. G. G. M. 2005, *The Physics and Chemistry of the Interstellar Medium* (*The Physics and Chemistry of the Interstellar Medium*, by A. G. G. M. Tielens, pp. . ISBN 0521826349. Cambridge, UK: Cambridge University Press, 2005.)
- Ulich, B. L. & Haas, R. W. 1976, *ApJS*, 30, 247
- van Broekhuizen, F. A., Keane, J. V., & Schutte, W. A. 2004, *A&A*, 415, 425
- van der Tak, F. F. S. 2005, in *IAU Symposium*, Vol. 227, *Massive Star Birth: A Crossroads of Astrophysics*, ed. R. Cesaroni, M. Felli, E. Churchwell, & M. Walmsley, 70–79
- Wakelam, V., Caselli, P., Ceccarelli, C., Herbst, E., & Castets, A. 2004, *A&A*, 422, 159
- Wakelam, V., Hersant, F., & Herpin, F. 2011, *A&A*, 529, A112+
- Werner, M. W., Becklin, E. E., Gatley, I., et al. 1979, *MNRAS*, 188, 463
- Whittet, D. C. B., Schutte, W. A., Tielens, A. G. G. M., et al. 1996, *A&A*, 315, L357
- Willner, S. P., Gillett, F. C., Herter, T. L., et al. 1982, *ApJ*, 253, 174
- Wilson, T. L. 1999, *Reports on Progress in Physics*, 62, 143
- Zinnecker, H. & Yorke, H. W. 2007, *ARA&A*, 45, 481